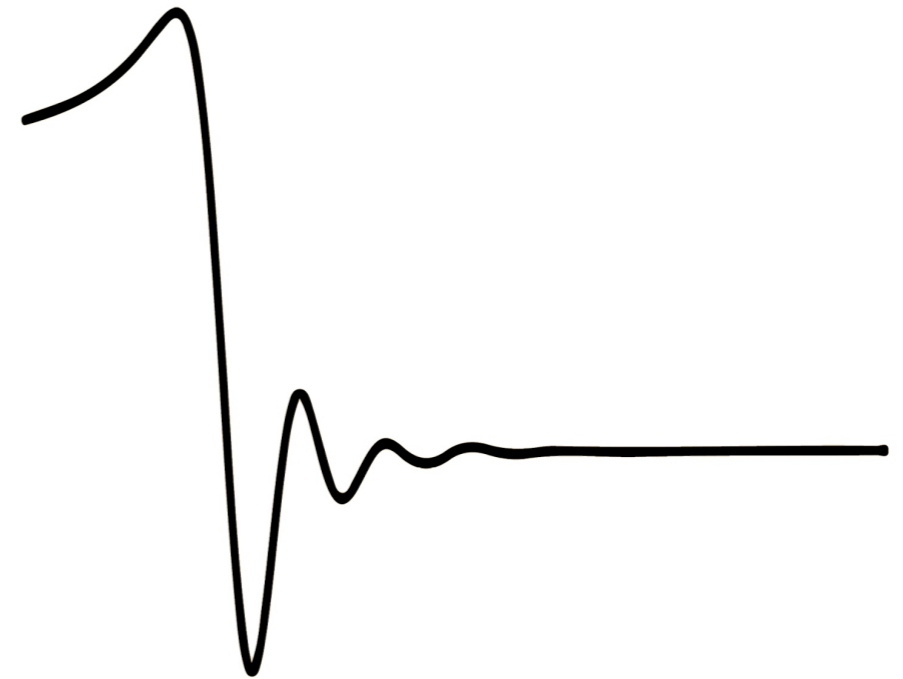


Astrophysical black holes: theory and observations

@ 59th Winter School “Gravity: Classical, Quantum and Phenomenology”

Wroclaw, 12-21 February 2023



Andrea Maselli

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BH Sociology



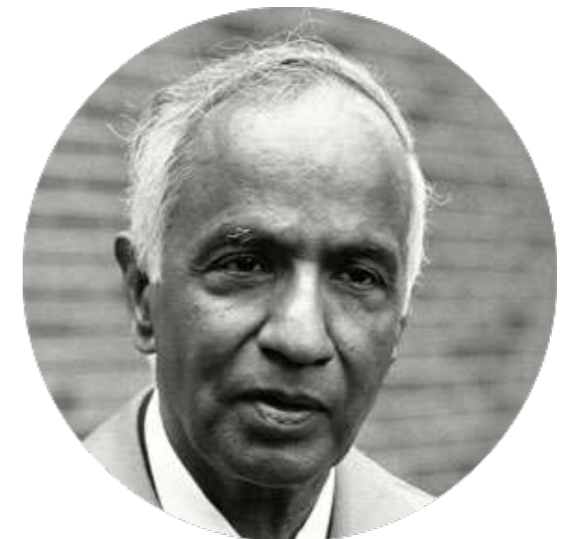
*“As you see, the **war** treated me kindly enough, in spite of the heavy gunfire, to allow me to get away from it all and take this walk in the land of your **ideas**: ...”*

Karl Schwarzschild to Albert Einstein
Letter dated 22 December 1915

*“In my entire scientific life, extending over forty-five years, the most shattering experience has been the realisation that an **exact** solution of Einstein’s equations of general relativity provides the absolutely exact representation of untold numbers of black holes that populate the universe.”*

S. Chandrasekhar

The Nora and Edward Ryerson lecture, Chicago April 22 1975



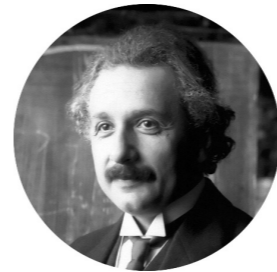
History of BHs



Mitchell



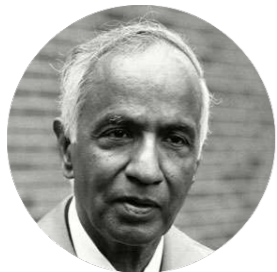
Schwarzschild



Einstein



Eddington



Chandrasekhar



Kerr



Wheeler



Oppenheimer



Penrose



Carter



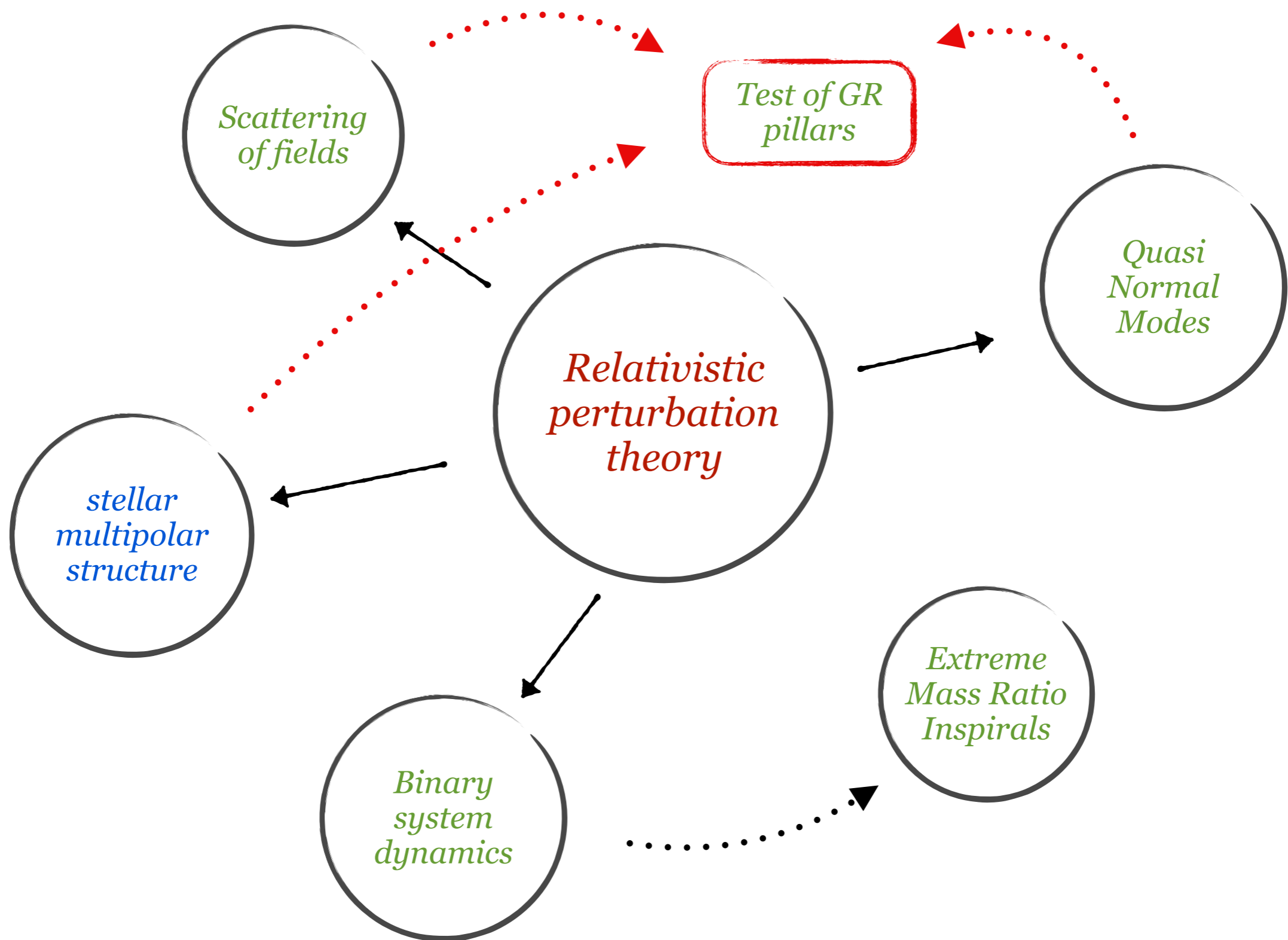
Hawking



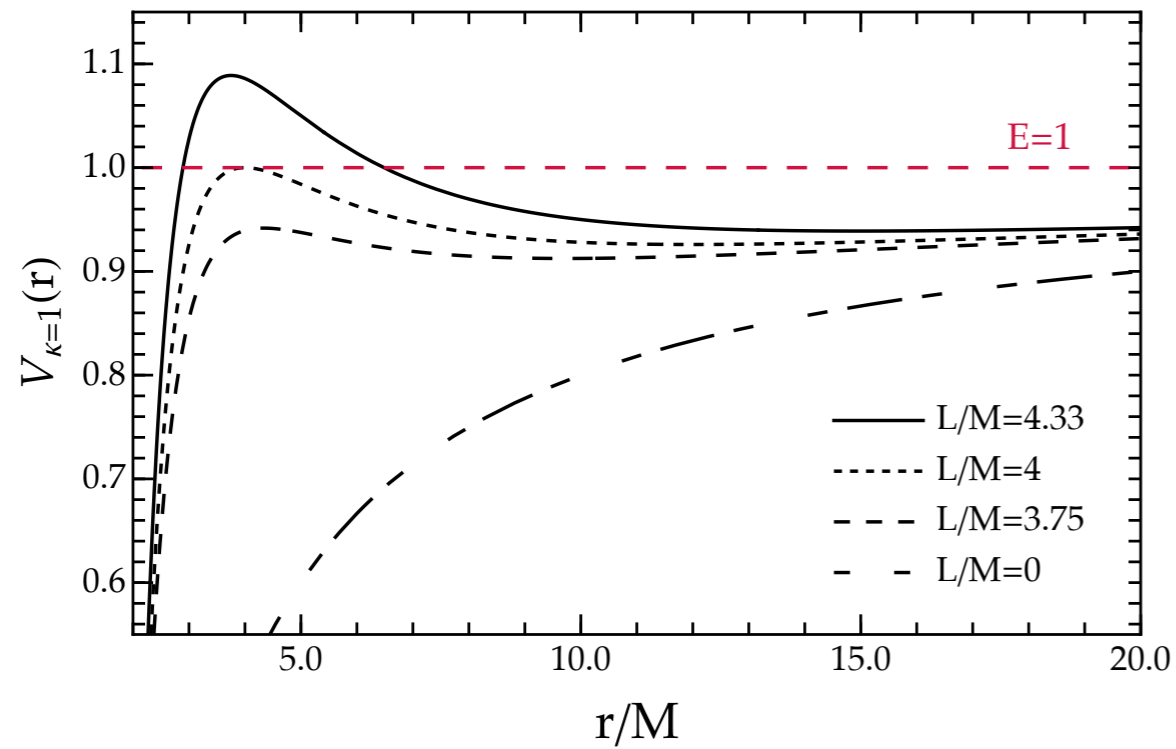
Thorne

Many more...

BH perturbations

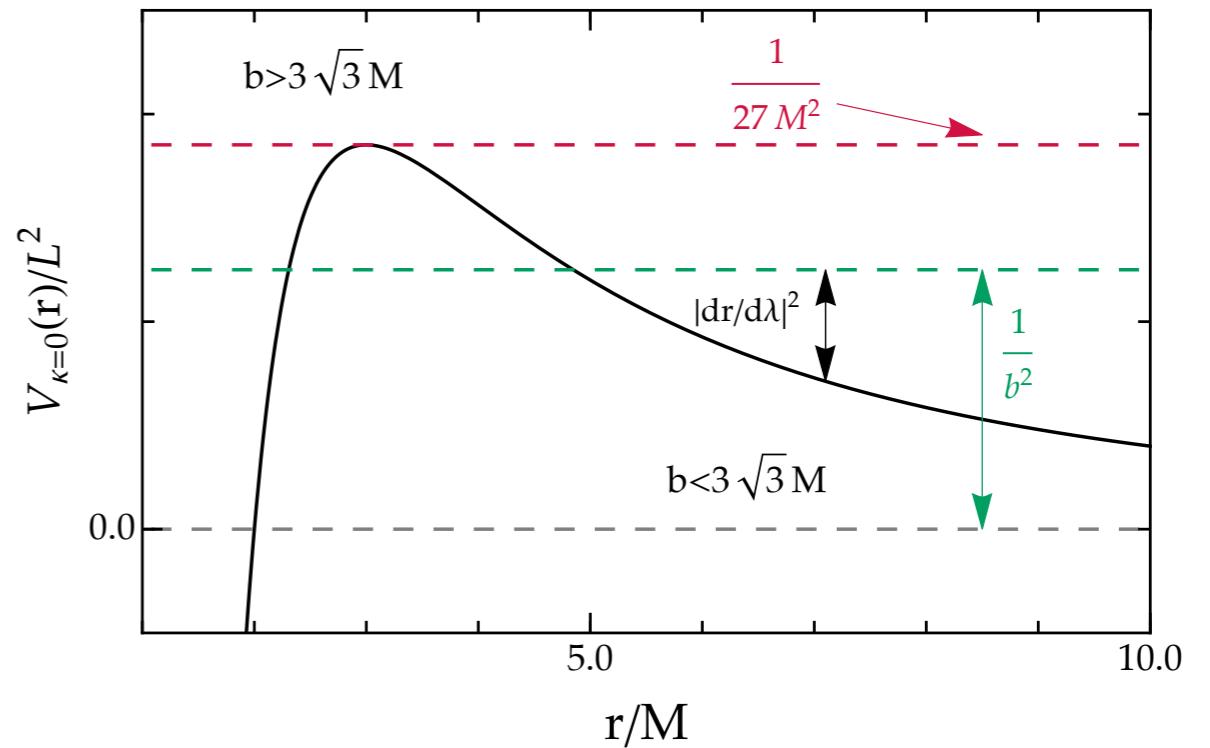


Effective potentials



← massive particles

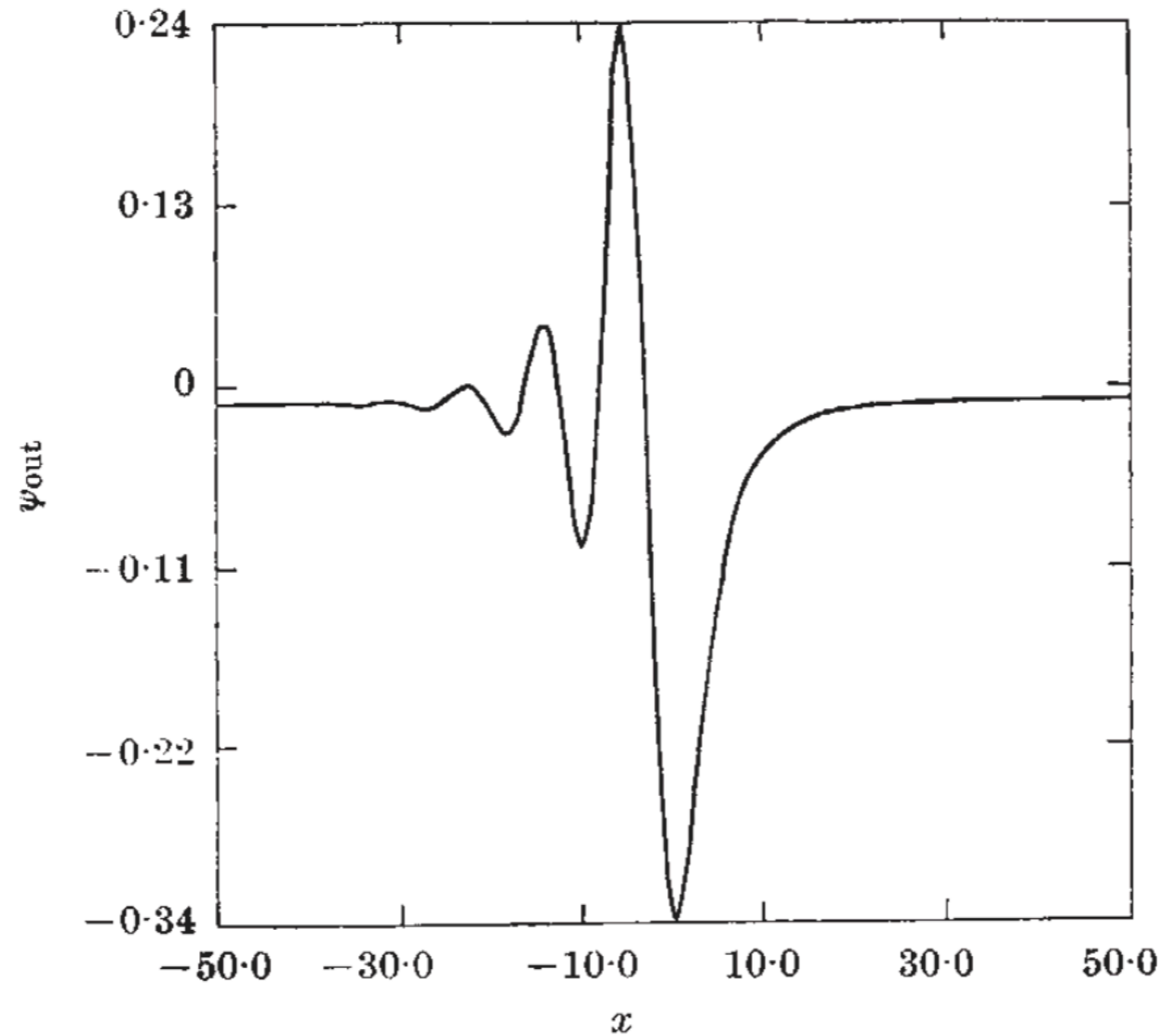
massless particles →



QNM universal behaviour

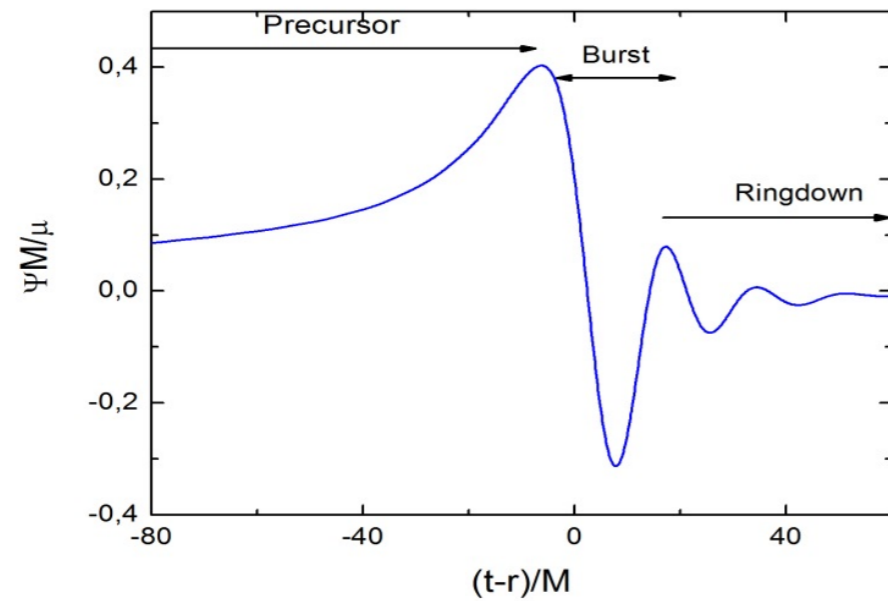
Scattering experiments which show always the same behaviour

C. V. Vishveshwara, Nature 227 (1970)

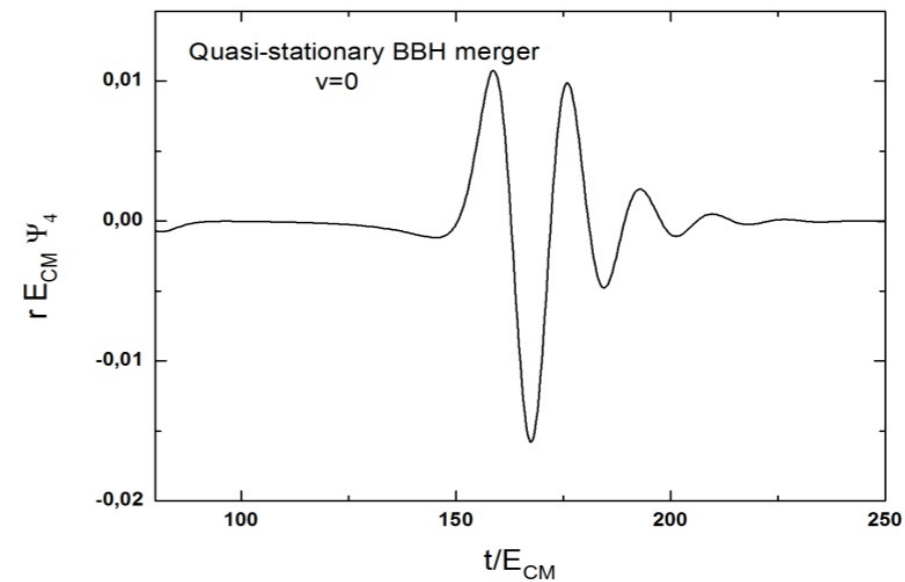


○ *Scattering of initial gaussian packages on Schwarzschild BH*

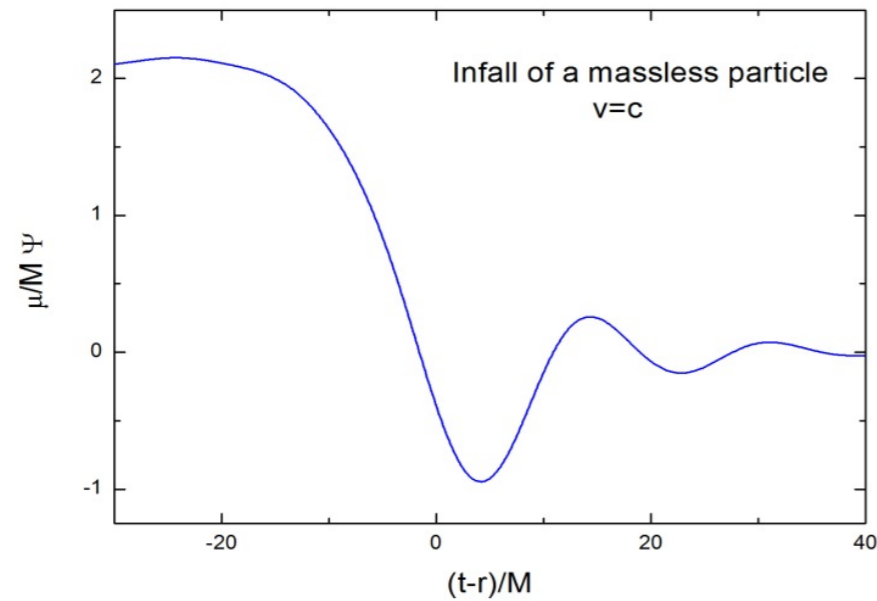
QNM universal behaviour



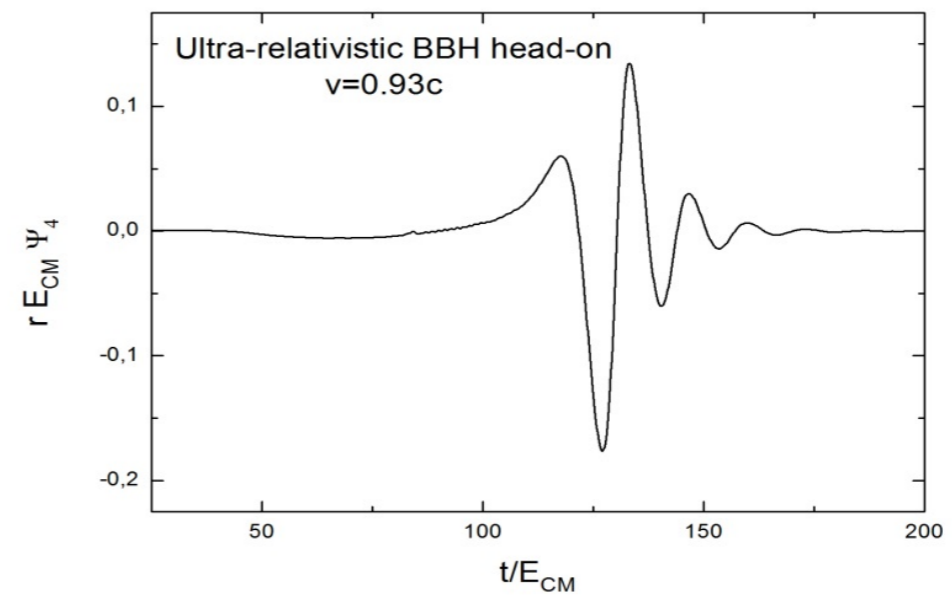
M. Davis +, Phys. Rev. Lett. 27 (1971)



P. Anninos +, Phys. Rev. Lett. 71 (1993)



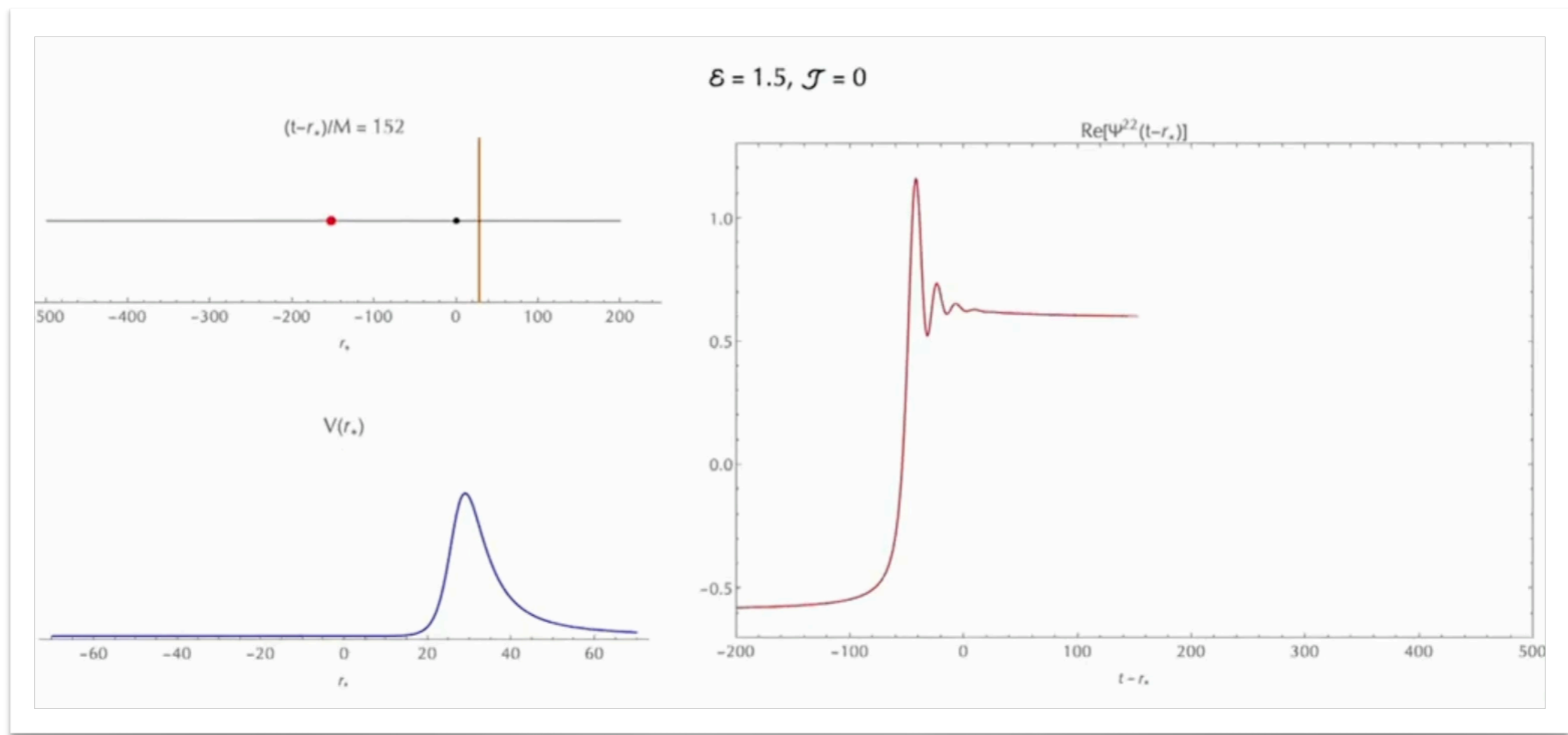
V. Cardoso & J. Lemos, Phys. Lett. B 538 (2002)



U. Sperhake +, Phys. Rev. Lett. 101, 161101 (2008)

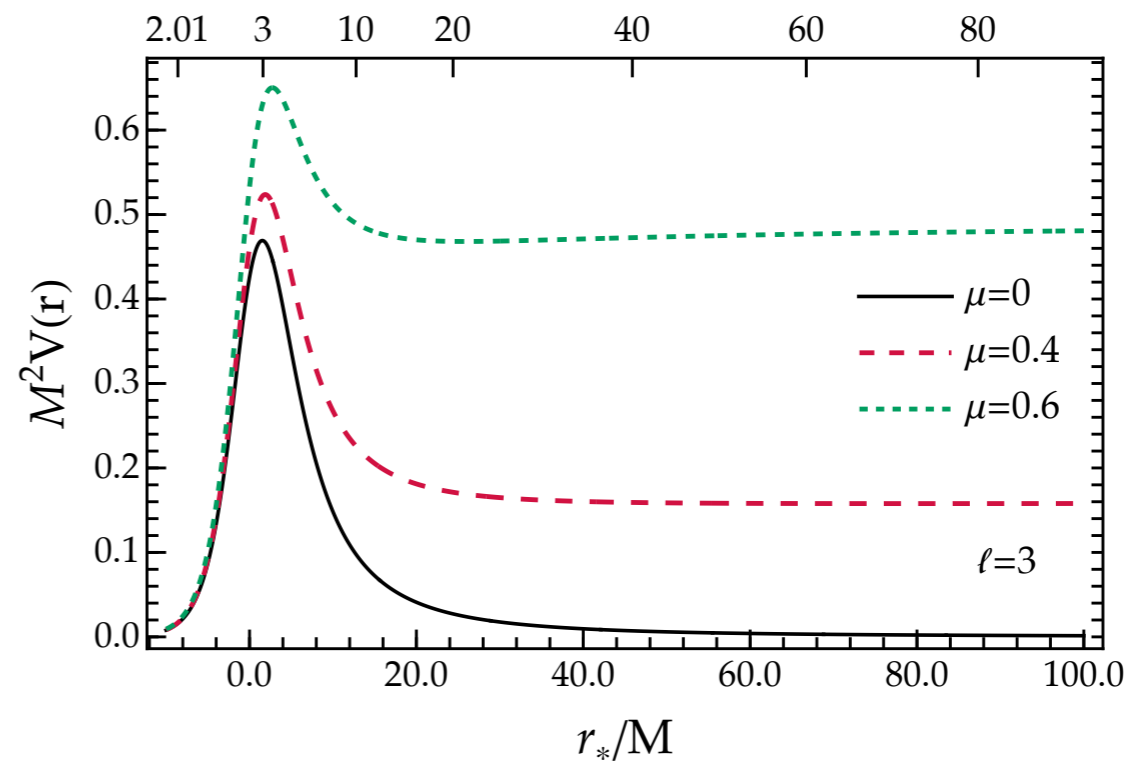
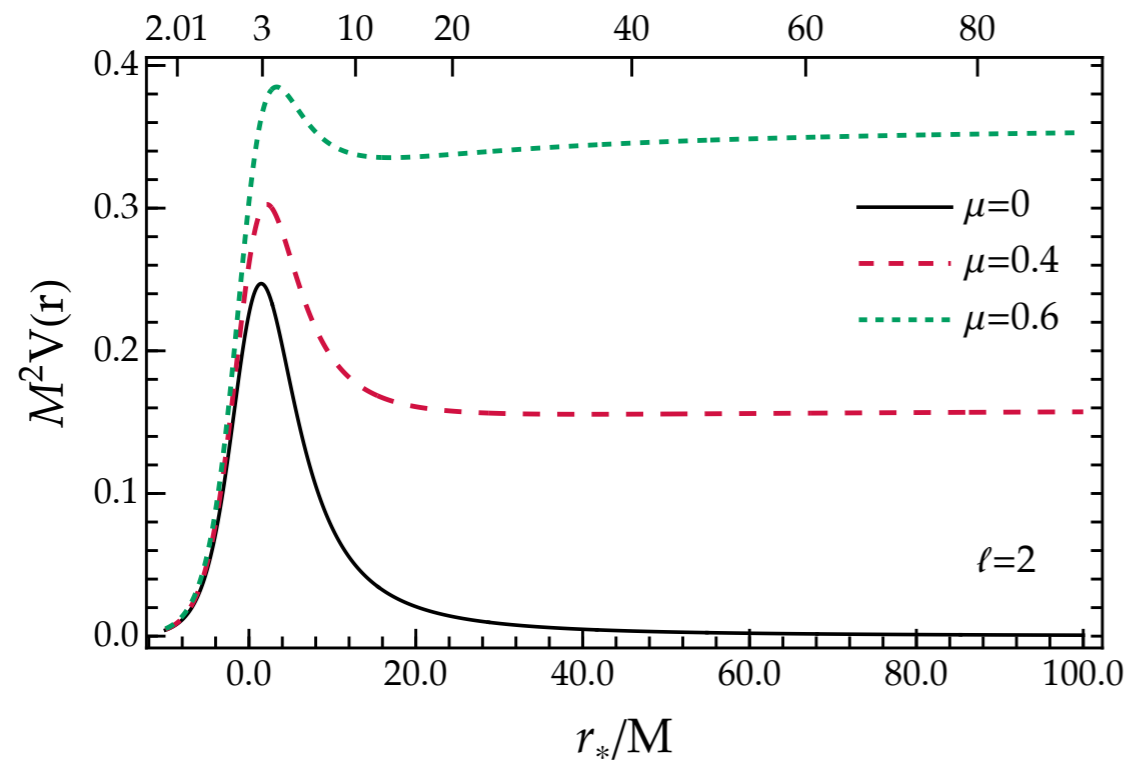
$$e^{-0.0898 t/M_{BH}} \sin(0.374 t/M_{BH})$$

QNMs and the light ring



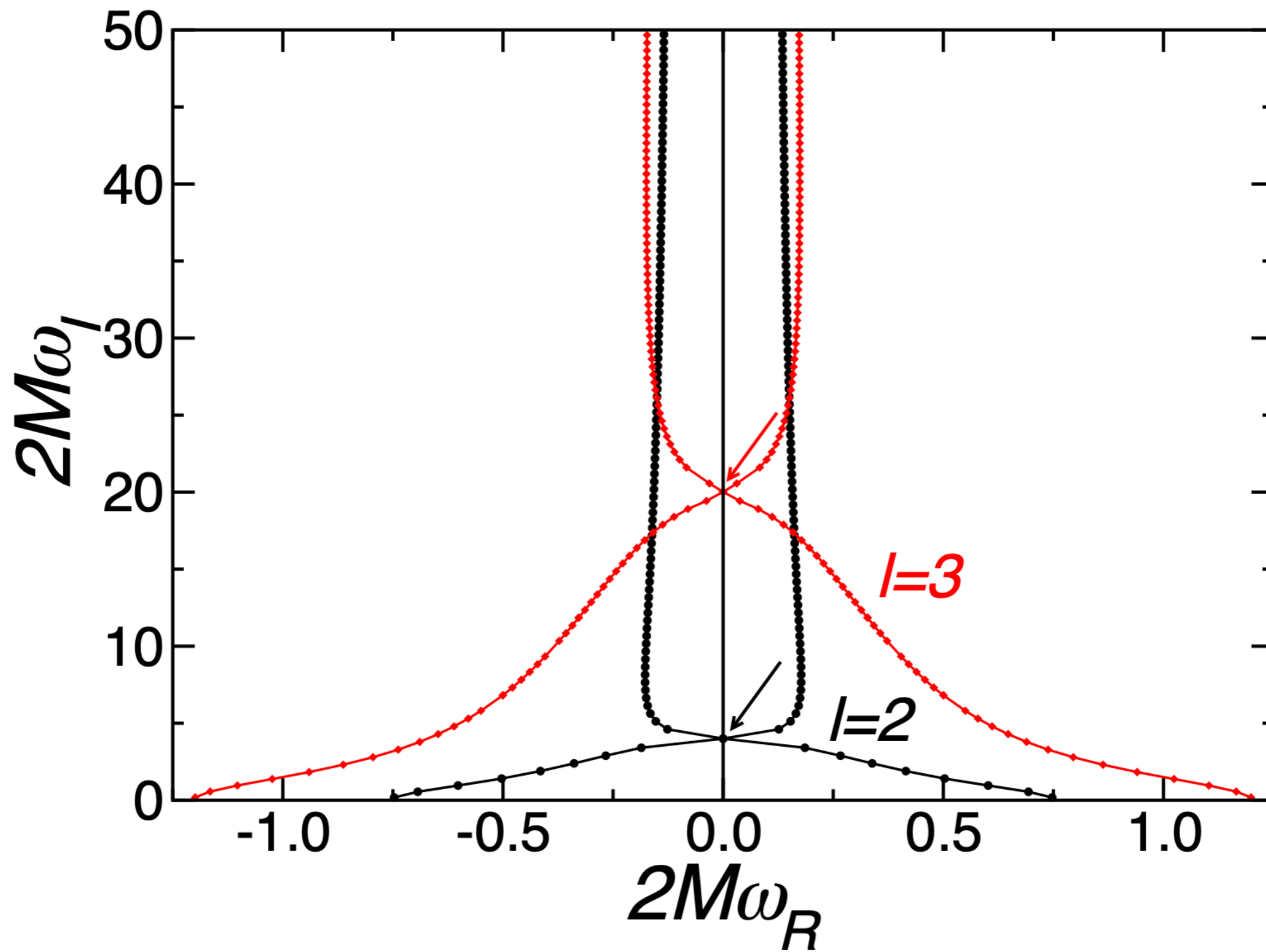
Scattering potentials

Scattering potentials for scalar field perturbations



Schwarzschild frequencies

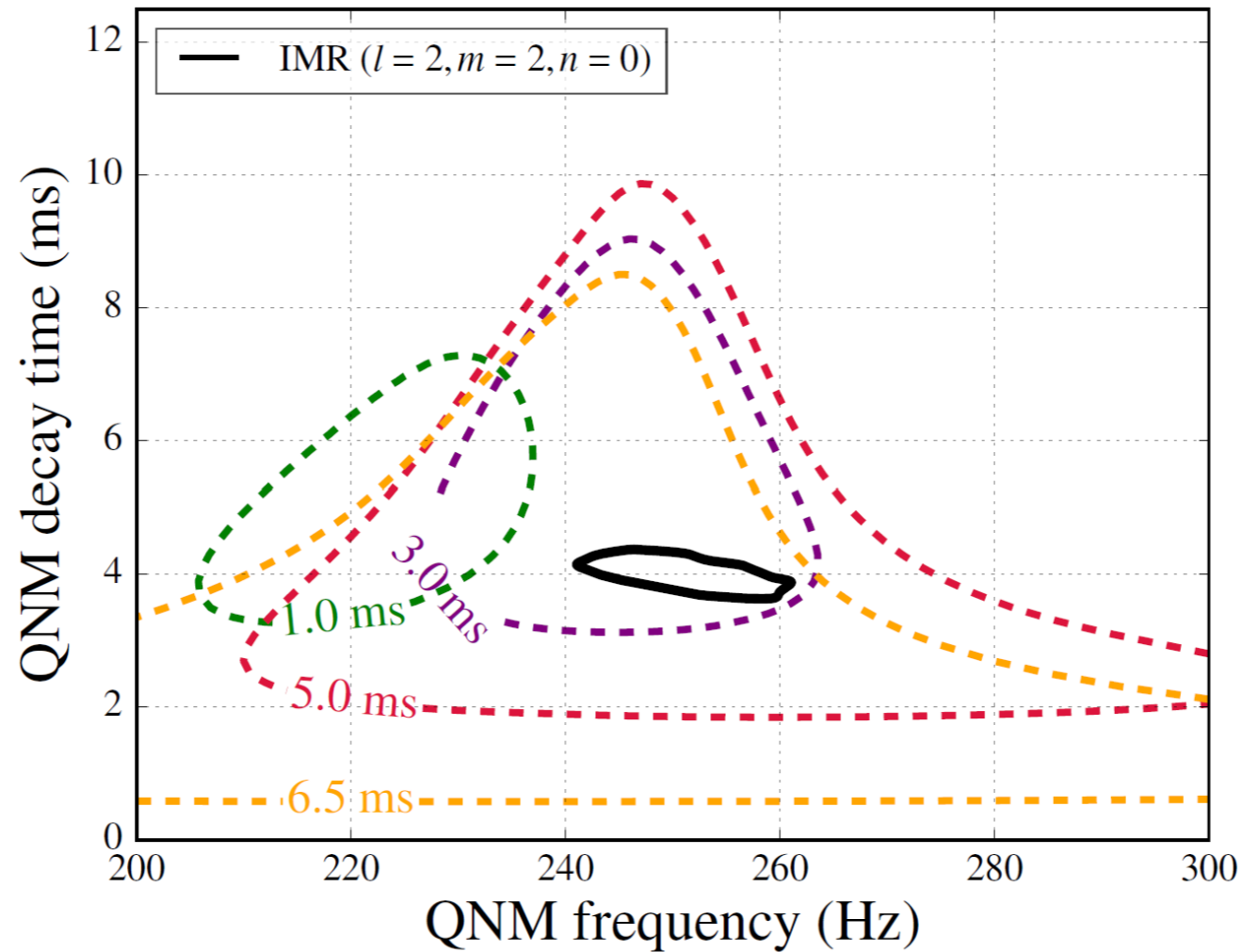
Credit Emanuele Berti



The start of BH spectroscopy

90% posterior distributions on the QNM frequencies from GW150914

LIGO/Virgo, Phys. Rev. Lett. 116, 221101 (2016)



○ Black solid is 90% posterior of QNM as derived from the posterior mass and spin of remnant

The start of BH spectroscopy

Frequencies and damping times for more events

LIGO/Virgo, Phys. Rev. D 103, 122002 (2021)

Event	Redshifted frequency [Hz]			Redshifted damping time [ms]		
	IMR	DS	pSEOB	IMR	DS	pSEOB
GW150914	248 ⁺⁸ ₋₇	247 ⁺¹⁴ ₋₁₆	—	4.2 ^{+0.3} _{-0.2}	4.8 ^{+3.7} _{-1.9}	—
GW170104	287 ⁺¹⁵ ₋₂₅	228 ⁺⁷¹ ₋₁₀₂	—	3.5 ^{+0.4} _{-0.3}	3.6 ^{+36.2} _{-2.1}	—
GW170814	293 ⁺¹¹ ₋₁₄	527 ⁺³⁴⁰ ₋₃₃₂	—	3.7 ^{+0.3} _{-0.2}	25.1 ^{+22.2} _{-19.0}	—
GW170823	197 ⁺¹⁷ ₋₁₇	222 ⁺⁶⁶⁴ ₋₆₂	—	5.5 ^{+1.0} _{-0.8}	13.4 ^{+31.8} _{-9.8}	—
GW190408_181802	319 ⁺¹¹ ₋₂₀	504 ⁺⁴⁷⁹ ₋₄₅₉	—	3.2 ^{+0.3} _{-0.3}	10.0 ^{+32.5} _{-8.9}	—
GW190421_213856	162 ⁺¹⁴ ₋₁₄	—	171 ⁺⁵⁰ ₋₁₆	6.3 ^{+1.2} _{-0.8}	—	8.5 ^{+5.3} _{-4.2}
GW190503_185404	191 ⁺¹⁷ ₋₁₅	—	265 ⁺⁵⁰¹ ₋₇₉	5.3 ^{+0.8} _{-0.8}	—	3.5 ^{+3.4} _{-1.8}
GW190512_180714	381 ⁺³³ ₋₄₂	220 ⁺⁶⁸⁶ ₋₄₂	—	2.6 ^{+0.2} _{-0.2}	26.1 ^{+21.3} _{-22.9}	—
GW190513_205428	241 ⁺²⁶ ₋₂₈	250 ⁺⁴⁹³ ₋₈₈	—	4.3 ^{+1.1} _{-0.4}	5.3 ^{+19.2} _{-3.8}	—
GW190519_153544	127 ⁺⁹ ₋₉	123 ⁺¹¹ ₋₁₉	124 ⁺¹² ₋₁₃	9.5 ^{+1.7} _{-1.5}	9.7 ^{+9.0} _{-3.8}	10.3 ^{+3.6} _{-3.1}
GW190521	68 ⁺⁴ ₋₄	65 ⁺³ ₋₃	67 ⁺² ₋₂	15.8 ^{+3.9} _{-2.5}	22.1 ^{+12.4} _{-7.4}	30.7 ^{+7.7} _{-7.4}
GW190521_074359	198 ⁺⁷ ₋₇	197 ⁺¹⁵ ₋₁₅	205 ⁺¹⁵ ₋₁₂	5.4 ^{+0.4} _{-0.4}	7.7 ^{+6.4} _{-3.3}	5.3 ^{+1.5} _{-1.2}
GW190602_175927	105 ⁺¹⁰ ₋₉	93 ⁺¹³ ₋₂₂	99 ⁺¹⁵ ₋₁₅	10.0 ^{+2.0} _{-1.4}	10.0 ^{+17.2} _{-4.5}	8.8 ^{+5.4} _{-3.6}
GW190706_222641	108 ⁺¹¹ ₋₁₀	109 ⁺⁷ ₋₁₂	112 ⁺⁷ ₋₈	10.9 ^{+2.4} _{-2.2}	20.4 ^{+25.2} _{-12.9}	19.4 ^{+7.2} _{-8.9}
GW190708_232457	497 ⁺¹⁰ ₋₄₆	642 ⁺²⁷⁹ ₋₅₉₆	—	2.1 ^{+0.2} _{-0.1}	24.6 ^{+23.0} _{-22.6}	—
GW190727_060333	178 ⁺¹⁸ ₋₁₆	345 ⁺⁵⁸⁷ ₋₂₆₇	201 ⁺¹¹ ₋₂₁	6.1 ^{+1.1} _{-0.8}	21.1 ^{+25.6} _{-17.9}	15.4 ^{+5.3} _{-6.1}
GW190828_063405	239 ⁺¹⁰ ₋₁₁	247 ⁺³⁵⁰ ₋₁₅	—	4.8 ^{+0.6} _{-0.5}	17.3 ^{+25.3} _{-10.4}	—
GW190910_112807	177 ⁺⁸ ₋₈	166 ⁺⁹ ₋₈	174 ⁺¹² ₋₈	5.9 ^{+0.8} _{-0.5}	13.2 ^{+17.1} _{-6.2}	9.5 ^{+3.1} _{-2.7}
GW190915_235702	232 ⁺¹⁴ ₋₁₈	534 ⁺³⁷¹ ₋₄₉₃	—	4.6 ^{+0.8} _{-0.6}	15.0 ^{+30.1} _{-13.1}	—

The start of BH spectroscopy

New tests are available once QNM spectroscopy is in full power

LIGO/Virgo, Phys. Rev. D 103, 122002 (2021)

- Masses and spins of the final object produced by the merger

Event	Redshifted final mass $(1+z)M_f [M_\odot]$				Final spin χ_f			
	IMR	Kerr ₂₂₀	Kerr ₂₂₁	Kerr _{HM}	IMR	Kerr ₂₂₀	Kerr ₂₂₁	Kerr _{HM}
	GW150914	68.8 ^{+3.6} _{-3.1}	62.7 ^{+19.0} _{-12.1}	71.7 ^{+13.2} _{-12.5}	80.3 ^{+20.1} _{-21.7}	0.69 ^{+0.05} _{-0.04}	0.52 ^{+0.33} _{-0.44}	0.69 ^{+0.18} _{-0.36}
GW170104	58.5 ^{+4.6} _{-4.1}	56.2 ^{+19.1} _{-11.6}	61.3 ^{+16.7} _{-13.2}	104.3 ^{+207.7} _{-43.1}	0.66 ^{+0.08} _{-0.11}	0.26 ^{+0.42} _{-0.24}	0.51 ^{+0.34} _{-0.44}	0.59 ^{+0.34} _{-0.51}
GW170814	59.7 ^{+3.0} _{-2.3}	46.1 ^{+133.0} _{-33.6}	56.6 ^{+20.9} _{-11.1}	171.2 ^{+268.7} _{-143.5}	0.72 ^{+0.07} _{-0.05}	0.52 ^{+0.42} _{-0.47}	0.47 ^{+0.40} _{-0.42}	0.54 ^{+0.41} _{-0.48}
GW170823	88.8 ^{+11.2} _{-10.2}	73.8 ^{+26.8} _{-23.7}	79.0 ^{+21.3} _{-13.2}	103.0 ^{+133.1} _{-46.7}	0.72 ^{+0.09} _{-0.12}	0.46 ^{+0.40} _{-0.41}	0.36 ^{+0.38} _{-0.32}	0.74 ^{+0.22} _{-0.61}
GW190408_181802	53.0 ^{+3.2} _{-3.4}	22.4 ^{+253.0} _{-11.1}	46.6 ^{+18.8} _{-10.9}	127.4 ^{+327.7} _{-107.6}	0.67 ^{+0.06} _{-0.07}	0.45 ^{+0.45} _{-0.40}	0.36 ^{+0.46} _{-0.33}	0.46 ^{+0.47} _{-0.41}
GW190512_180714	43.5 ^{+4.0} _{-2.8}	37.6 ^{+48.9} _{-22.4}	36.7 ^{+19.3} _{-24.8}	99.4 ^{+247.6} _{-66.5}	0.65 ^{+0.07} _{-0.07}	0.41 ^{+0.47} _{-0.37}	0.45 ^{+0.40} _{-0.39}	0.77 ^{+0.20} _{-0.66}
GW190513_205428	70.6 ^{+11.5} _{-6.7}	55.5 ^{+31.5} _{-42.1}	68.5 ^{+28.2} _{-11.8}	88.7 ^{+250.0} _{-41.9}	0.68 ^{+0.14} _{-0.12}	0.38 ^{+0.48} _{-0.34}	0.31 ^{+0.53} _{-0.28}	0.59 ^{+0.34} _{-0.52}
GW190519_153544	146.8 ^{+14.7} _{-15.4}	120.7 ^{+39.7} _{-21.5}	125.9 ^{+24.3} _{-21.7}	155.4 ^{+84.4} _{-42.5}	0.79 ^{+0.07} _{-0.13}	0.42 ^{+0.41} _{-0.36}	0.52 ^{+0.25} _{-0.40}	0.70 ^{+0.21} _{-0.50}
GW190521	256.6 ^{+36.6} _{-30.4}	282.2 ^{+50.0} _{-61.9}	284.0 ^{+40.4} _{-43.9}	299.3 ^{+57.7} _{-62.4}	0.71 ^{+0.12} _{-0.16}	0.76 ^{+0.14} _{-0.38}	0.78 ^{+0.10} _{-0.22}	0.80 ^{+0.13} _{-0.30}
GW190521_074359	88.0 ^{+4.3} _{-4.8}	83.0 ^{+24.0} _{-17.2}	86.4 ^{+14.1} _{-14.8}	105.9 ^{+20.8} _{-26.4}	0.72 ^{+0.05} _{-0.07}	0.57 ^{+0.31} _{-0.49}	0.67 ^{+0.17} _{-0.34}	0.87 ^{+0.09} _{-0.39}
GW190602_175927	163.8 ^{+20.7} _{-18.3}	156.4 ^{+71.4} _{-30.6}	160.0 ^{+37.4} _{-31.2}	261.7 ^{+84.4} _{-91.5}	0.70 ^{+0.10} _{-0.14}	0.34 ^{+0.41} _{-0.31}	0.46 ^{+0.31} _{-0.39}	0.79 ^{+0.14} _{-0.49}
GW190706_222641	171.1 ^{+20.0} _{-23.7}	136.0 ^{+52.0} _{-29.3}	152.5 ^{+37.8} _{-28.4}	184.0 ^{+139.2} _{-55.8}	0.78 ^{+0.09} _{-0.18}	0.41 ^{+0.42} _{-0.37}	0.55 ^{+0.31} _{-0.45}	0.68 ^{+0.26} _{-0.54}
GW190708_232457	34.4 ^{+2.7} _{-0.7}	28.9 ^{+285.4} _{-17.9}	32.3 ^{+15.0} _{-12.2}	171.9 ^{+307.6} _{-147.8}	0.69 ^{+0.04} _{-0.04}	0.47 ^{+0.45} _{-0.42}	0.34 ^{+0.44} _{-0.31}	0.43 ^{+0.51} _{-0.39}
GW190727_060333	99.2 ^{+10.7} _{-9.8}	78.7 ^{+45.7} _{-66.4}	88.8 ^{+25.7} _{-16.0}	107.4 ^{+112.1} _{-42.7}	0.73 ^{+0.10} _{-0.10}	0.53 ^{+0.42} _{-0.47}	0.45 ^{+0.39} _{-0.41}	0.71 ^{+0.24} _{-0.59}
GW190828_063405	75.7 ^{+6.0} _{-5.2}	71.2 ^{+35.8} _{-55.5}	69.6 ^{+22.0} _{-17.3}	99.0 ^{+166.0} _{-49.1}	0.75 ^{+0.06} _{-0.07}	0.72 ^{+0.25} _{-0.62}	0.65 ^{+0.27} _{-0.55}	0.92 ^{+0.06} _{-0.74}
GW190910_112807	97.0 ^{+9.3} _{-7.1}	112.2 ^{+32.0} _{-31.7}	107.7 ^{+28.6} _{-27.4}	137.1 ^{+59.5} _{-31.4}	0.70 ^{+0.08} _{-0.07}	0.76 ^{+0.18} _{-0.55}	0.75 ^{+0.17} _{-0.46}	0.91 ^{+0.07} _{-0.27}
GW190915_235702	74.8 ^{+7.9} _{-7.4}	38.3 ^{+335.1} _{-27.4}	63.0 ^{+19.1} _{-9.9}	137.3 ^{+324.1} _{-96.2}	0.70 ^{+0.09} _{-0.11}	0.52 ^{+0.43} _{-0.46}	0.27 ^{+0.40} _{-0.24}	0.55 ^{+0.39} _{-0.49}

as predicted from the
inspiral

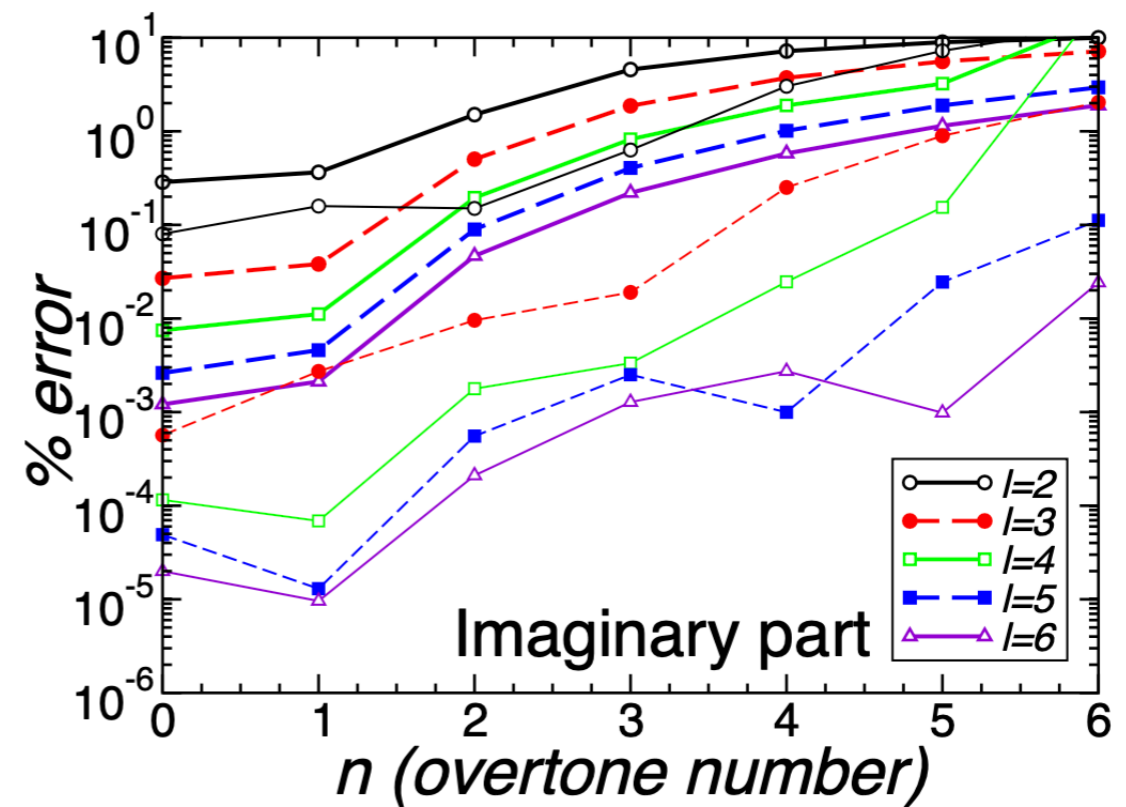
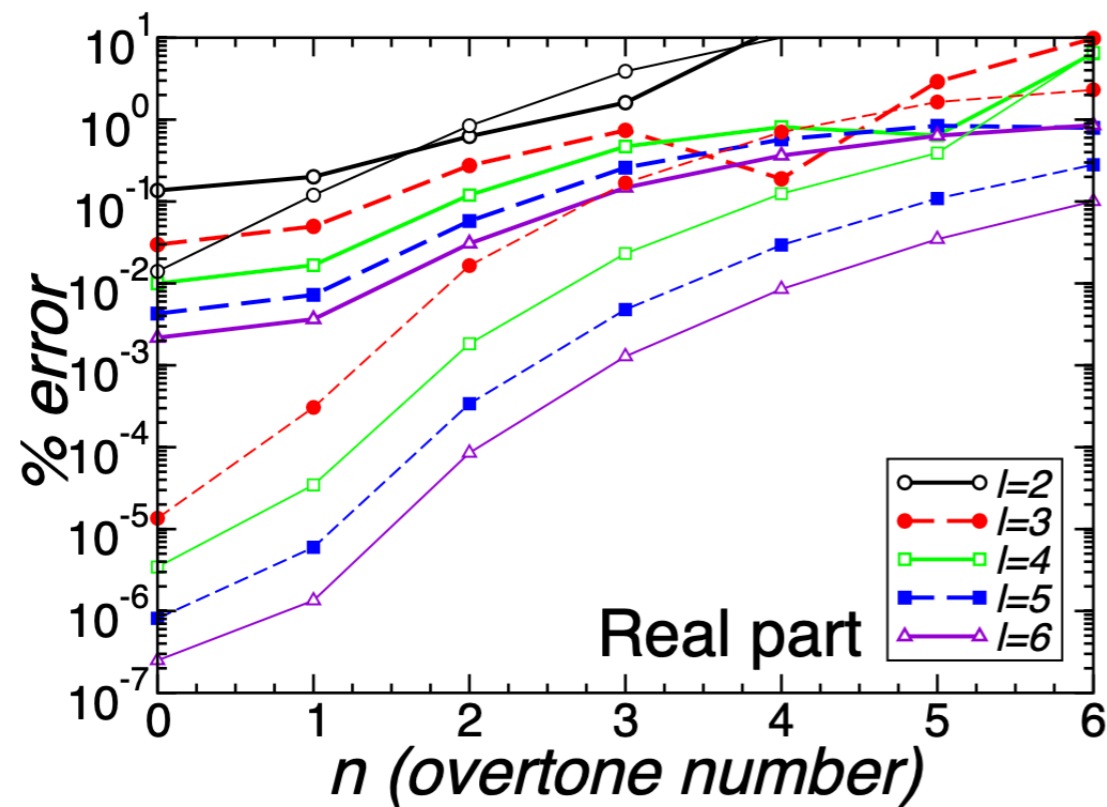


from the post-merger only

WKB vs “exact” values

Relative error on the QNM frequencies of a Schwarzschild BH, computed through WKB and continued fraction

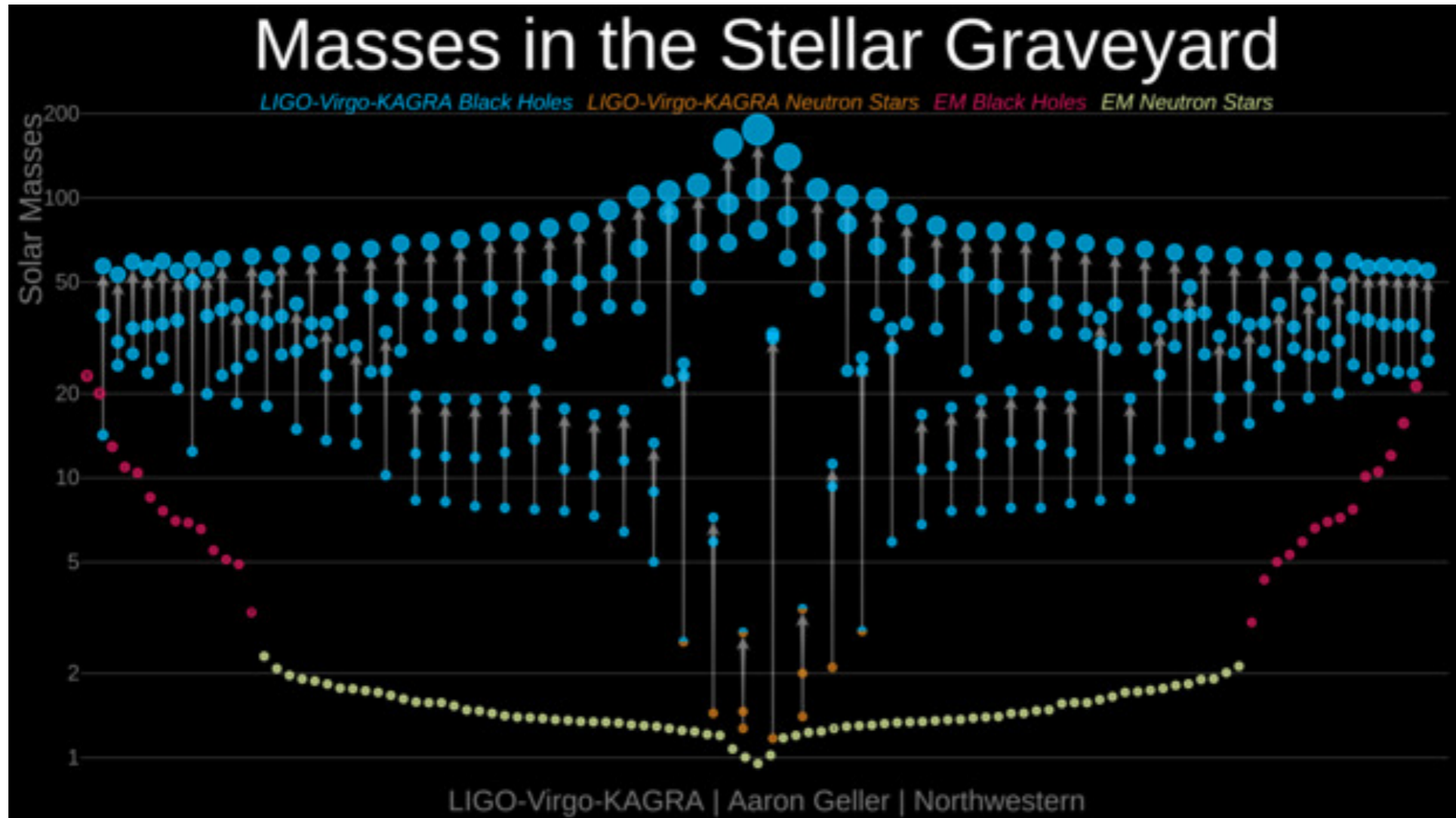
E. Berti +, Class. Quant. Grav. 26, 163001 (2010)



Exotic Love

Alone in the dark?

Are Kerr black hole alone in the Universe?



Alone in the dark?

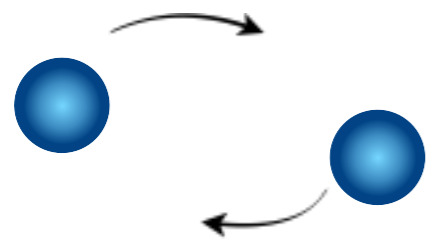
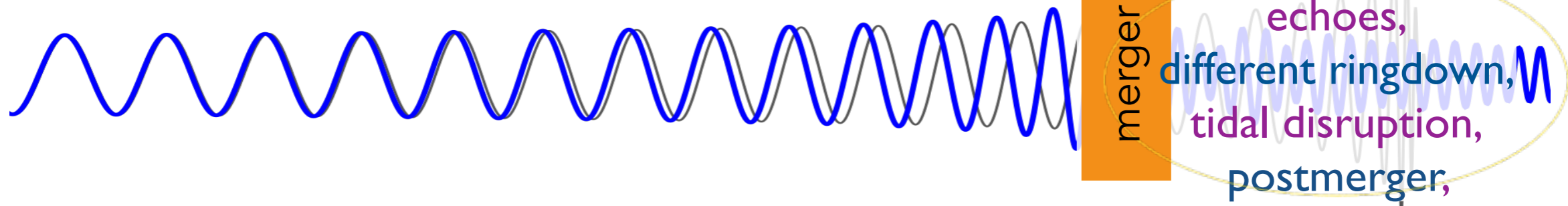
Kerr BHs as endpoint of stellar evolution?

- *Unexpected processes may avoid their formation*
- *Extended theories of gravity in which extra fields couple to the gravity sector can predict BHs different from Kerr solutions, with specific **hairs***
- *Other **Exotic Compact Objects** may be the output of stellar collapse, which form without an event horizon*
- *Can we distinguish ECOs with no horizon, compact enough to mimic a BH, **and/or**, BHs with different hairs?*

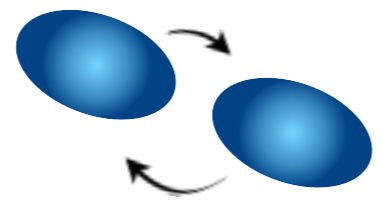
GW from coalescing binaries may provide new answers

What do we look for?

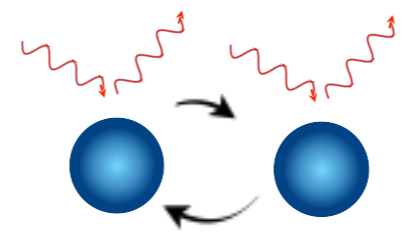
— BH-
BH — other
 objects



~point masses:
same signal
for all objects



tidal effects
+
spins
deformations



absence of horizon
absorption
effects



different ringdown
or
echoes

The Love number

Tidal interactions leave the footprint of the NS structure on the GW signal

*T. Hinderer, Astrophys. J. 677 (2008)
Damour & Nagar, Phys. Rev. D 80, 084035 (2009)
Binnington & Poisson, Phys. Rev. D 80, 084018 (2009)*

○ Deformation properties encoded within the *Love numbers*

The diagram consists of a central equation with two arrows pointing outwards from its top. The left arrow points to the text "star's quadrupole" and the right arrow points to the text "external tidal field".

$$Q_{ij} = \frac{2}{3} \kappa_2 R^5 \mathcal{E}_{ij} = \lambda \mathcal{E}_{ij}$$

star's quadrupole

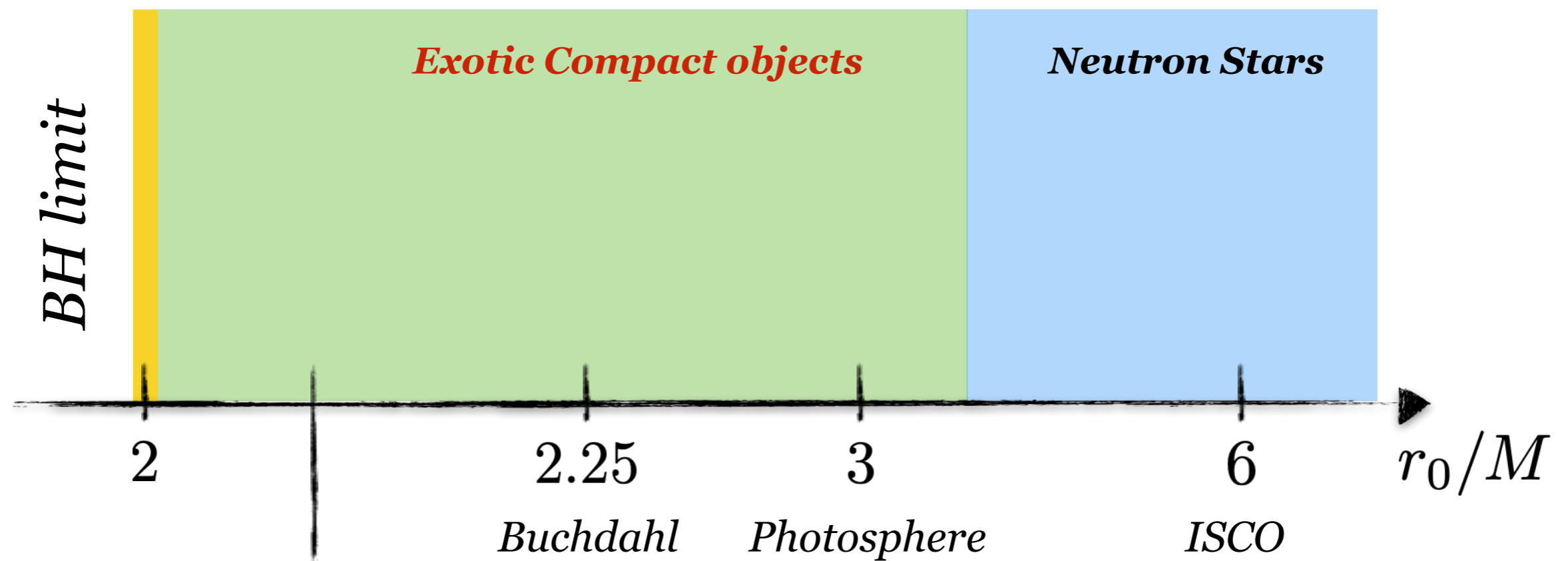
external tidal field

- λ depends on the internal structure only, for a given compactness
- λ enters within the gravitational waveform

ECOs vs BHs

A first order classification

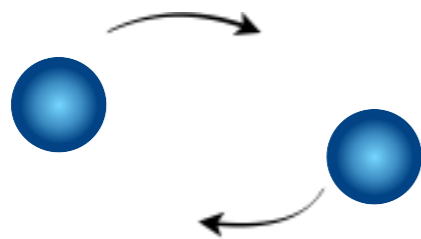
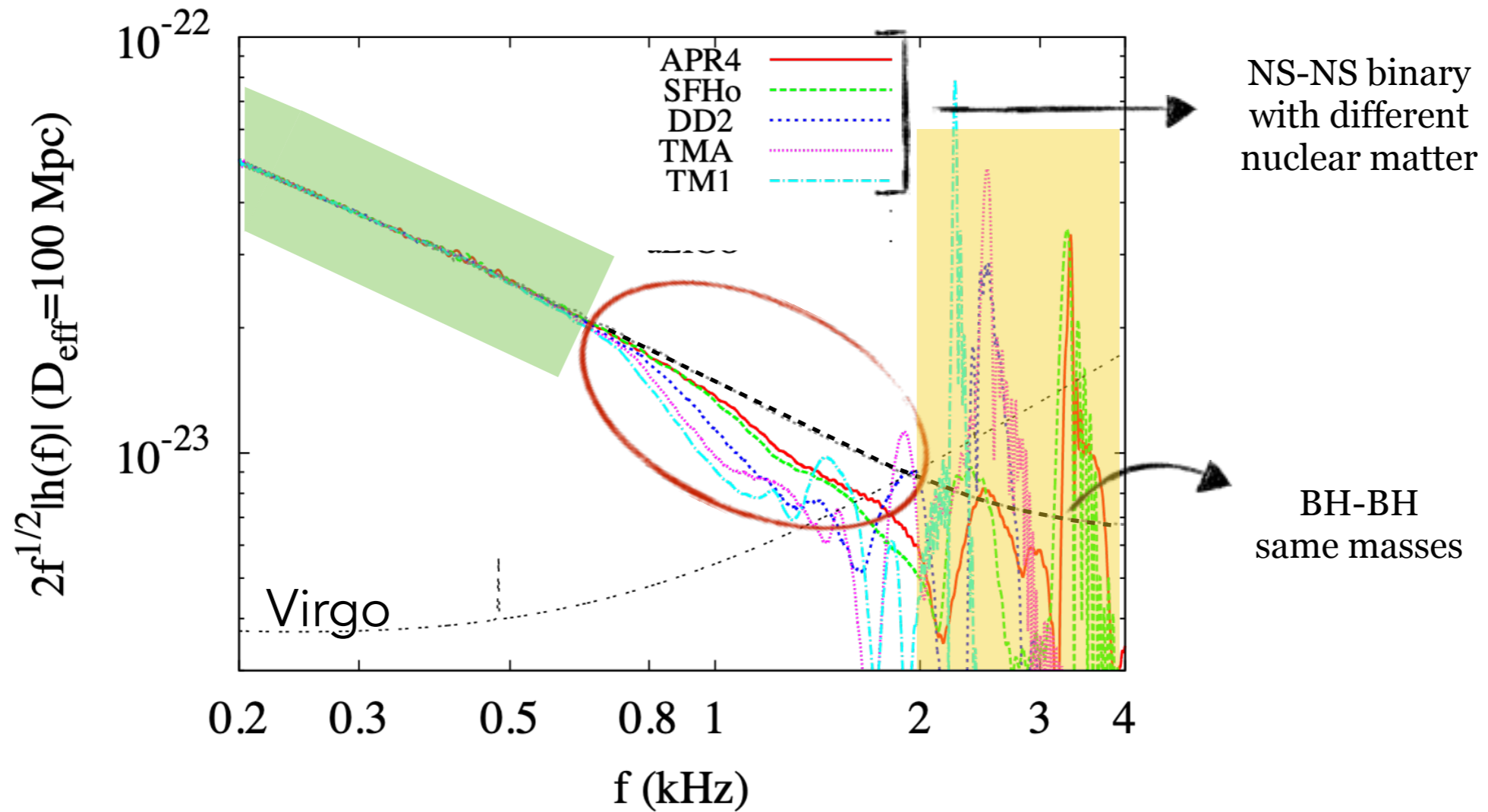
Cardoso & Pani, Living Rev.Rel. 22, 4 (2019)



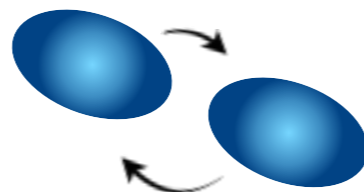
$$M_{BH}/R_{BH} = 0.5 \quad v.s. \quad M_{ECO}/R_{ECO} = 0.49(99\dots)$$

What do we look for?

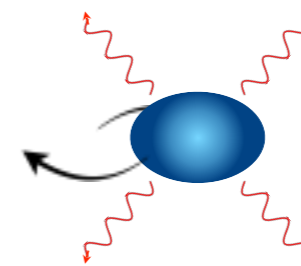
K. Hotokezaka +, Phys. Rev. D 93, 064082 (2016)



point masses



tidal interactions



merger/post-merger

The recipe of Love

Polar-electric-type perturbation of background metric

*Regge & Wheeler, PRD 108, 1063 (1957)
Zerilli, PRD 2, 2141 (1970)*

$$g_{\mu\nu} = g_{\mu\nu}^{(0)} + h_{\mu\nu}$$

$$\begin{pmatrix} -e^{\nu(r)} & 0 & 0 & 0 \\ 0 & e^{\lambda(r)} & 0 & 0 \\ 0 & 0 & r^2 & 0 \\ 0 & 0 & 0 & r^2 \sin^2 \theta \end{pmatrix} + \begin{pmatrix} -e^{\nu(r)} H_0(r) & 0 & 0 & 0 \\ 0 & e^{\lambda(r)} H_2(r) & 0 & 0 \\ 0 & 0 & r^2 K(r) & 0 \\ 0 & 0 & 0 & r^2 K(r) \sin^2 \theta \end{pmatrix} Y_{lm}(\theta, \phi)$$

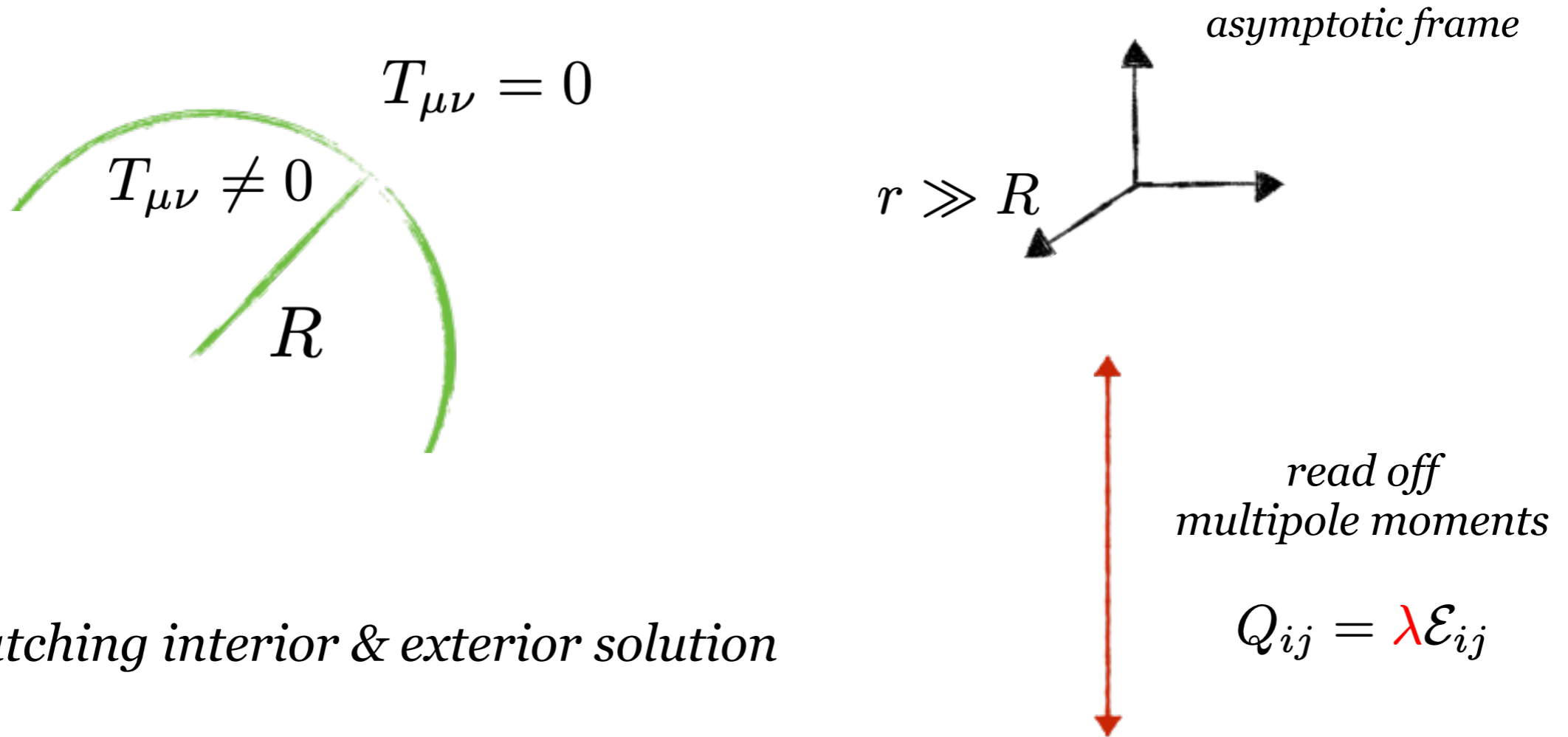
- *Solve at linear order in the perturbations H_0, H_2, K*
- *Cook everything within Einstein's equations $G_{\mu\nu} = kT_{\mu\nu}$*



set of sourced ODEs

The recipe of Love

- Numerically integrate from center to outside with appropriate boundary conditions



- Matching interior & exterior solution

$$g_{tt}^{(\text{int})} = g_{tt}^{(\text{ext})} = -1 + \frac{2M}{r} + \frac{3Q_{ij}\hat{n}^i\hat{n}^j}{r^3} - \mathcal{E}_{ij}x^ix^j$$

ECO Love numbers

Neutron stars

$$\mathcal{C} \in [0.1 \div 0.2]$$

$$\lambda \neq 0$$

GW170817



*constrain the NS
equation of state*

*Kerr Black
holes*

$$\mathcal{C} = 1/2$$

$$\lambda = 0$$

ECOs

$$\mathcal{C} < 0.5 \quad \lambda \neq 0$$

$$\mathcal{C} \rightarrow 0.5 \quad \lambda \rightarrow 0$$



*signature to distinguish
ECO and BH's inspiral*

*BH Beyond GR have non
vanishing Love number too!*

Landry & Poisson, *Phys. Rev. D* 91, 104018 (2015)
P. Pani +, *Phys. Rev. D* 92, 024010 (2017)
Le Tiec & Casals, *Phys. Rev. Lett.* 126, 131102 (2021)

V. Cardoso +, *Phys. Rev. D* 95, 084014 (2017)
A.M. +, *Phys. Rev. Lett.* 120, 081101 (2018)
A.M. +, *Class. Quant. Grav.* 36, 167001 (2010)

ECO Love numbers

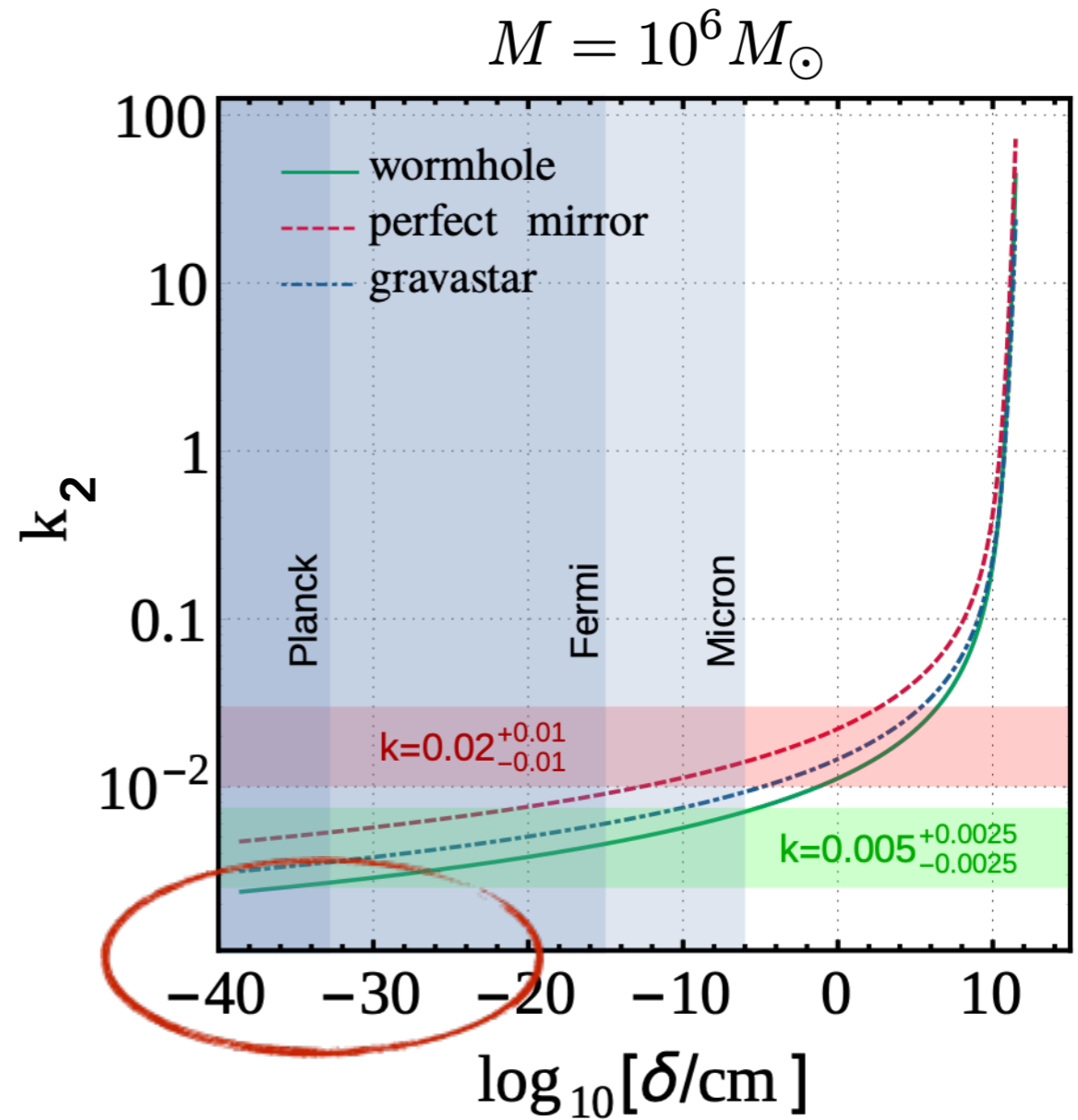
The love number reflects the distance of the ECO surface from its Schwarzschild radius

$$\delta \equiv r_0 - 2M \sim 2M e^{-1/k_2}$$

$$k_2 \simeq 0.005$$



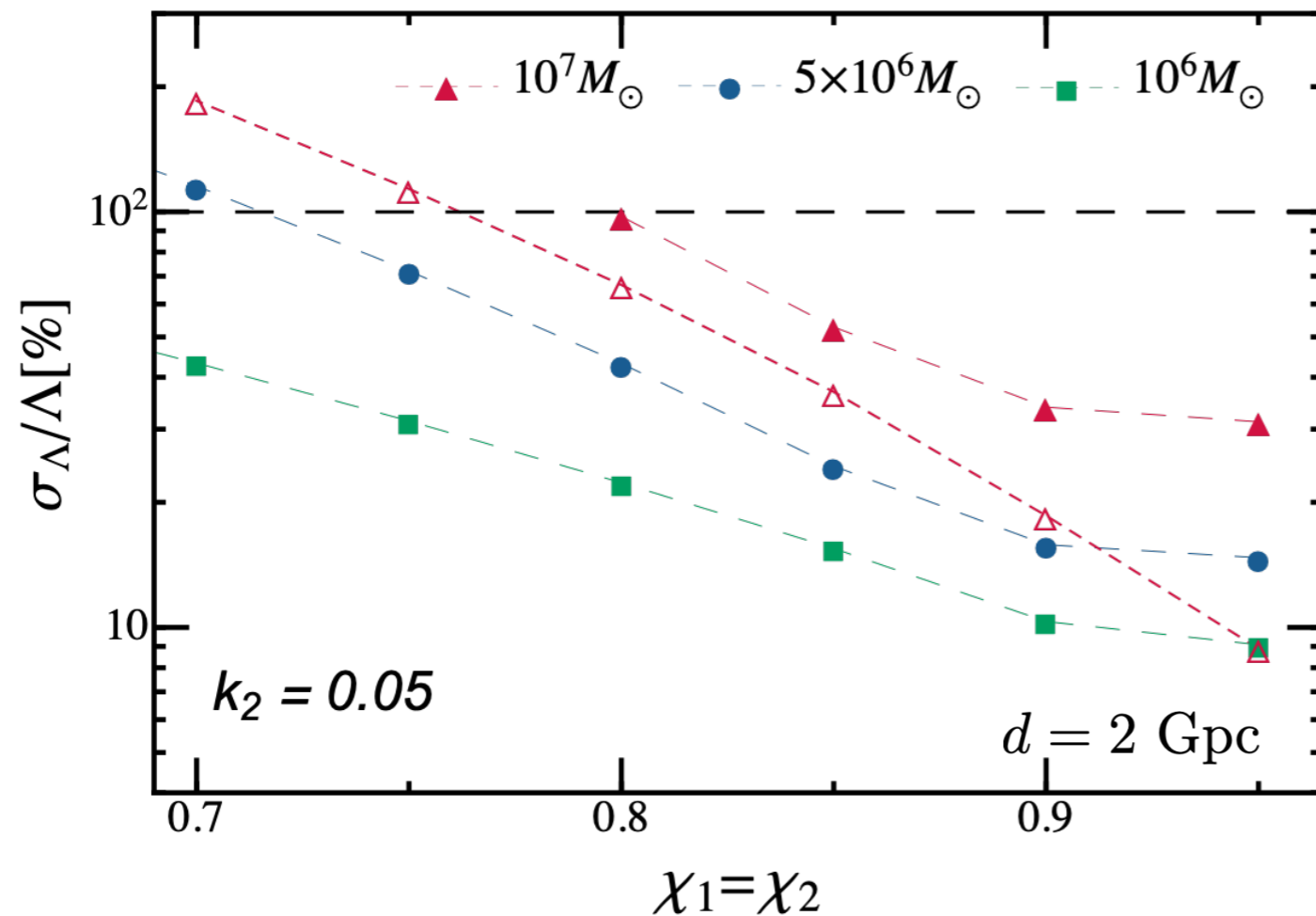
$$\delta \simeq 10^{-33} \text{ cm} \sim \ell_P$$



LISA & The uncertainty of Love

Spinning binaries can test microphysical modifications at the **horizon** scale

A.M. +, Phys. Rev. Lett. 120, 081101 (2018)



● $\sigma_{\Lambda}/\Lambda \sim 1/k_2$

$k_2 = 0.005$ @ $d = 2$ Gpc

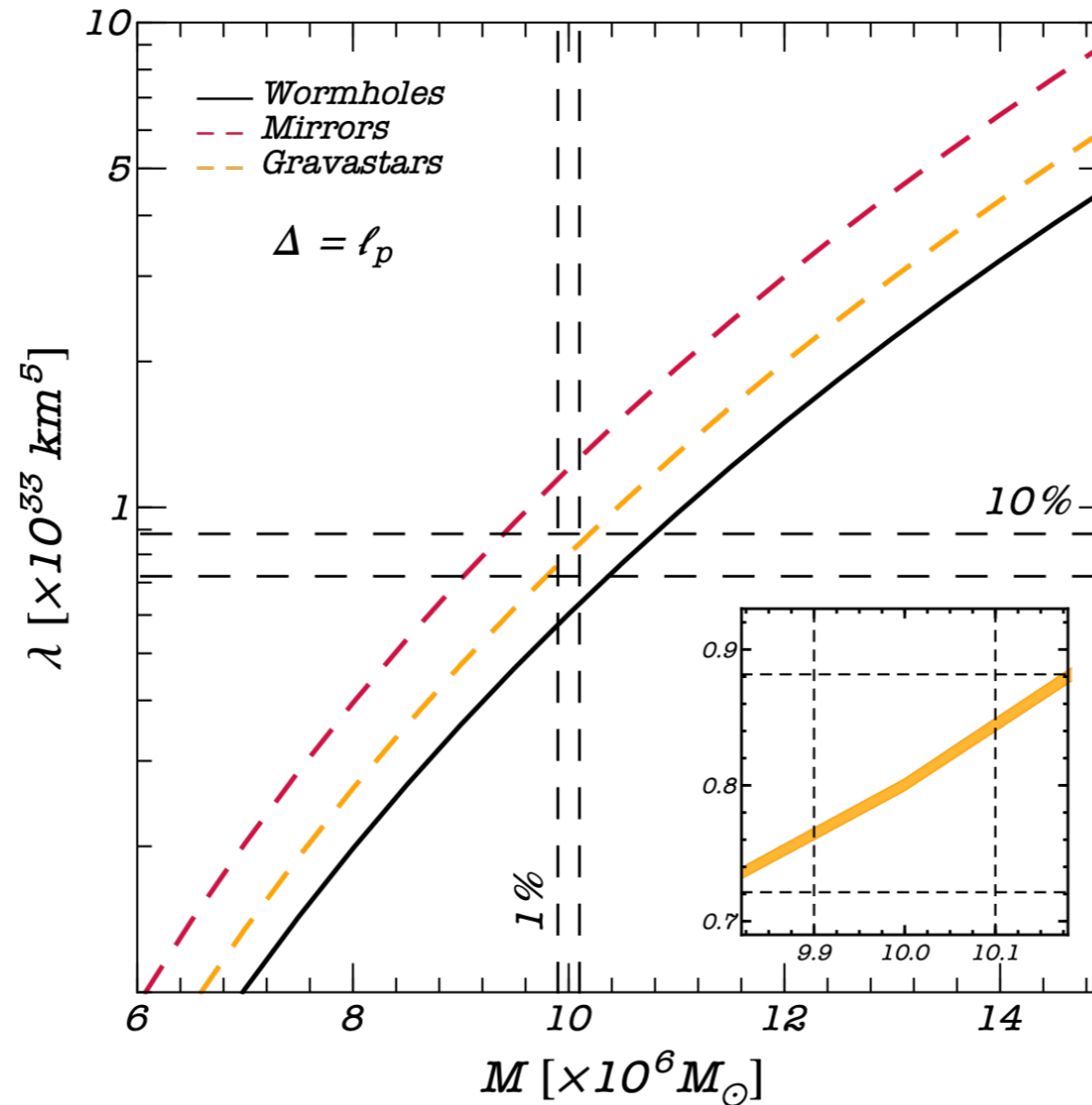
● $\sigma_{\Lambda}/\Lambda \sim d$

$k_2 = 0.025$ @ $d = 10$ Gpc

LISA & The uncertainty of Love

GWs can distinguish **between** models with quantum modifications

A.M. +, *Class. Quant. Grav.* 36, 167001 (2010)



Scalar fields

EMRI in nuce

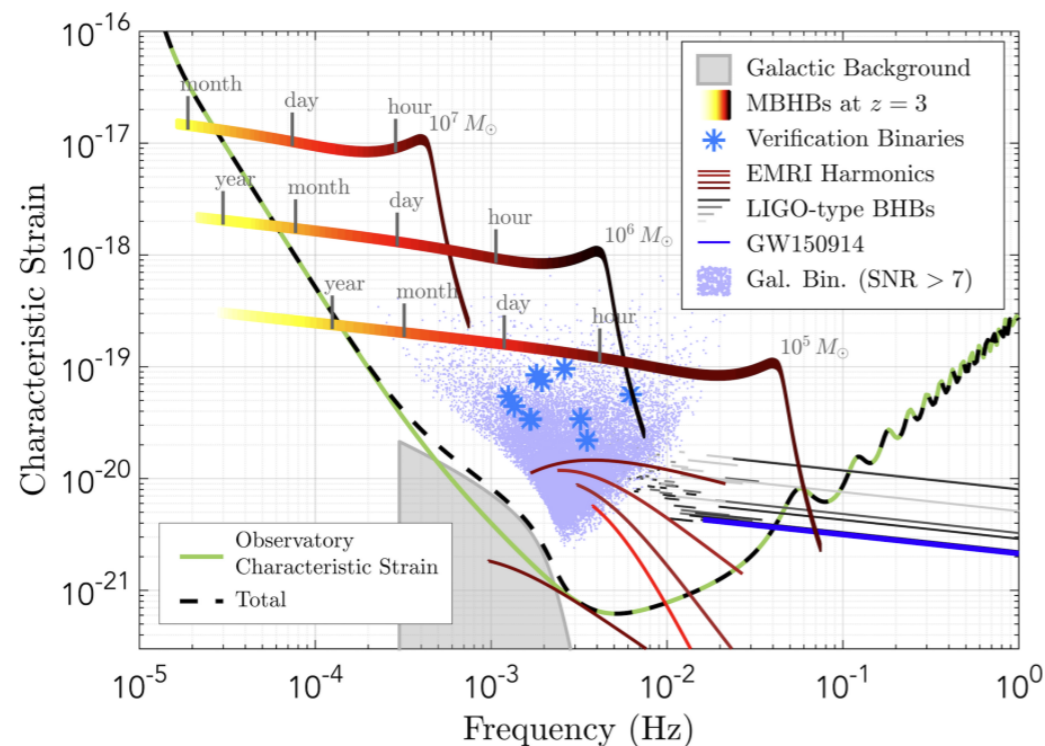
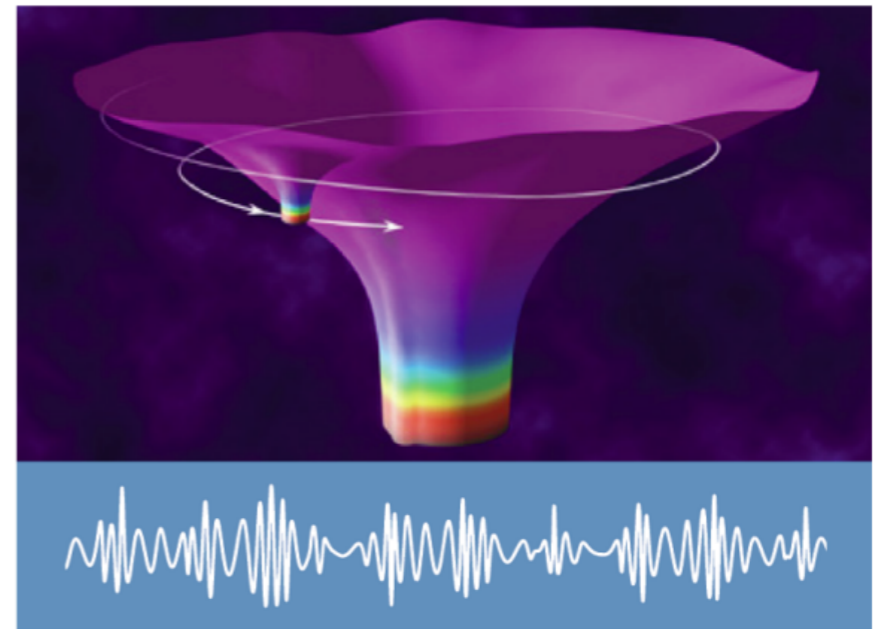
Binary systems with a stellar-mass body inspiralling into a massive black hole

○ *Primary with $M \sim (10^4 - 10^7)M_\odot$*

○ *Secondary such that the mass ratio*

$$q = m_p/M \sim (10^{-6} - 10^{-3})$$

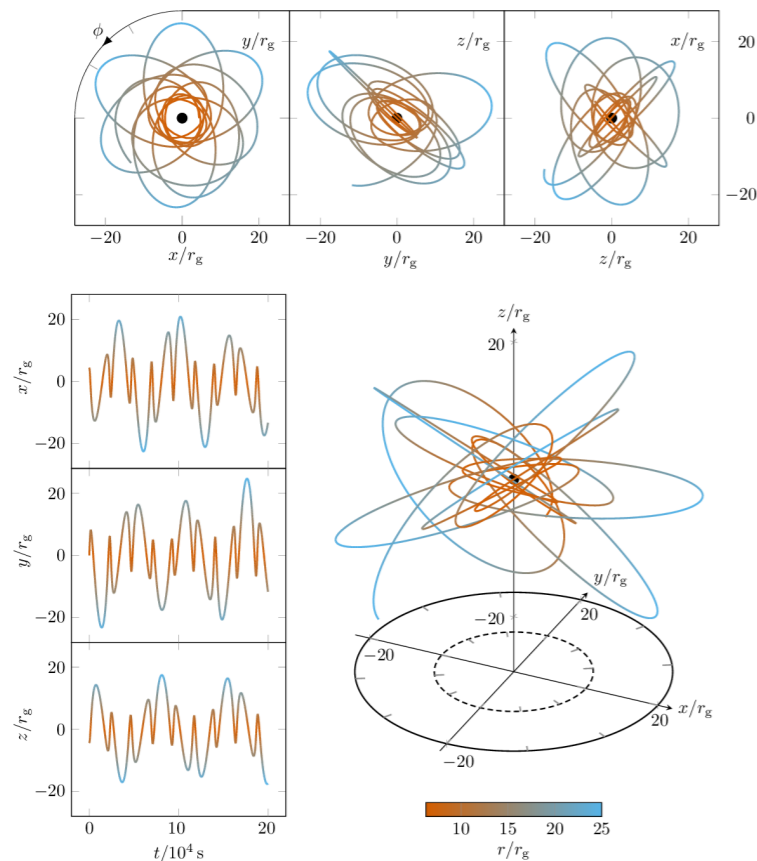
○ *Key point of theoretical description*



○ *Emit GWs in the mHz band, golden targets for LISA, dim to ground based detectors*

EMRI in nuce

EMRIs provide a rich phenomenology, due to their orbital features



Berry +, Astro2020 1903.03686 (2019)

- *Non equatorial orbits*
- *Eccentric motion*
- *Resonances*
- *Complete $\sim (10^4 - 10^5)$ cycles before the plunge*

bless and ***disguise***

Tracking EMRIs for $O(\text{year})$ requires accurate templates

Very appealing to test fundamental & astro-physics

Precise space-time map and accurate binary parameters

EMRI in GR


How do we study EMRI in GR?

- The asymmetric character introduces a natural parameter to study the problem in perturbation theory $q = m_p/M \ll 1$

$$g_{\alpha\beta} = g_{\alpha\beta}^0 + h_{\alpha\beta}$$

$$G_{\mu\nu} = T_{\mu\nu}^p = 8\pi m_p \int \frac{\delta^{(4)}(x - y_p(\lambda))}{\sqrt{-g}} \frac{dy_p^\alpha}{d\lambda} \frac{dy_p^\beta}{d\lambda} d\lambda$$

leading
adiabatic



Regge-Wheeler-Zerilli
(Schwarzschild)

Teukolsky
(Kerr)

- The solution determines the phase evolution

$$\phi(t) = \overset{\text{adiabatic}}{\phi_{\text{diss-1}}} + \overset{\text{first post-adiabatic}}{\dots}$$



$\mathcal{O}(1/q)$



$\mathcal{O}(1)$

The Setup

Scalar field φ non-minimally coupled to the gravity sectors

$$S[\mathbf{g}, \varphi, \Psi] = S_0[\mathbf{g}, \varphi] + \alpha S_c[\mathbf{g}, \varphi] + S_m[\mathbf{g}, \varphi, \Psi]$$

$$S_0 = \int d^4x \frac{\sqrt{-g}}{16\pi} \left(R - \frac{1}{2} \partial_\mu \varphi \partial^\mu \varphi \right)$$

Non-minimal
coupling

Matter fields

- Dimensionful coupling $[\alpha] = (\text{mass})^n$

We assume that

- BH solutions are connected to GR solutions $\alpha \rightarrow 0$
- S_c is analytic in φ

The Setup

Key simplifications for the **exterior** space-time occur for 2 families

1) Theories with no-hair theorems

2) Theories which evade no-hair but have dimensionful coupling α with $n \geq 1$

○ Any correction depend on $\zeta \equiv \frac{\alpha}{M^n} = q^n \frac{\alpha}{m_p^n} \ll 1$ $q = \frac{m_p}{M} \ll 1$

The exterior space-time can be approximated by the Kerr metric

For the **secondary**, consider the skeletonized approach

$$S_p = - \int m(\varphi) ds = - \int m(\varphi) \sqrt{g_{\mu\nu} \frac{dy_p^\mu}{d\lambda} \frac{dy_p^\nu}{d\lambda}} d\lambda$$

Eardley, ApJ 196 L59-62 (1975)
Damour & EF, CGQ 9, 9 (1992)

○ Extended body treated as point particle

○ $m(\varphi)$ scalar function

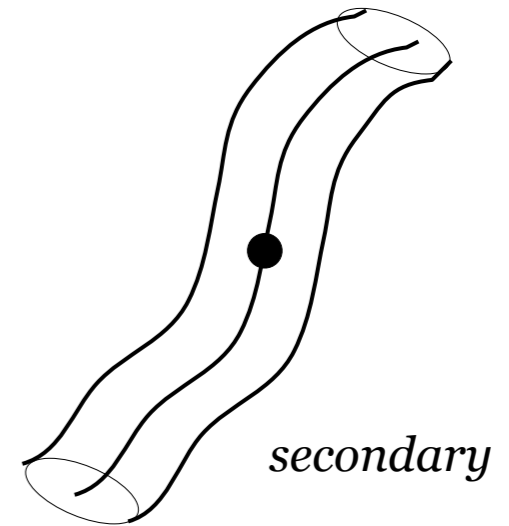
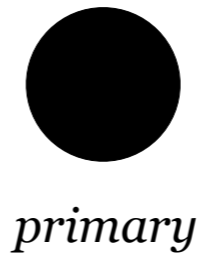
The Setup

Scale decoupling for the motion of the EMRI secondary

$l \ll L$

skeletonized
body

exterior
space-time



The orbital motion can be studied with perturbation theory in $q \ll 1$

- *GR modifications affect the motion of the particle but not the background*
- *The scalar field is a perturbation of a constant value $\varphi = \varphi_0 + \varphi_1$*

The field's equations

$(\zeta \sim q^n)$

In our units $[S_0] = (\text{mass})^2$ $[S_c] = (\text{mass})^{2-n}$ \longrightarrow $S_c \sim M^{-n} S_0$

$$G_{\mu\nu} = T_{\mu\nu}^{\text{scal}} + \alpha T_{\mu\nu}^c + T_{\mu\nu}^p$$

$$G_{\mu\nu} = \frac{1}{2} \cancel{\partial_\mu \varphi \partial_\nu \varphi} - \frac{1}{4} \cancel{g_{\mu\nu} (\partial\varphi)^2} + \cancel{-\alpha \frac{16\pi}{\sqrt{-g}} \frac{\delta S_c}{\delta g^{\mu\nu}}} + \int m(\varphi) \frac{\delta^{(4)}(x - y_p(\lambda))}{\sqrt{-g}} \frac{dy_p^\alpha}{d\lambda} \frac{dy_p^\beta}{d\lambda} d\lambda$$

quadratic in φ

$$\alpha T_{\mu\nu}^c \sim \zeta^2 G_{\mu\nu} \sim q^{2n} G_{\mu\nu}$$

$$\square\varphi + \cancel{\frac{8\pi\alpha}{\sqrt{-g}} \frac{\delta S_c}{\delta\varphi}} = 16\pi \int m'(\varphi) \frac{\delta^{(4)}(x - y_p(\lambda))}{\sqrt{-g}} d\lambda$$

$$\alpha \frac{\delta S_c}{\delta\varphi} \sim \zeta \square\varphi \ll \square\varphi$$

$m(\varphi), m'(\varphi)$ evaluated at the value of the exterior scalar field

Almost as in GR

From the scalar field equation inside the world tube, but far way to be weak field. In the body's frame

$$\varphi = \varphi_0 + \frac{m_p d}{\tilde{r}} + O\left(\frac{m_p^2}{\tilde{r}^2}\right) \quad \text{scalar charge}$$

○ *Matching with the scalar field equation outside the world tube*

$$m(\varphi_0) = m_p \quad \frac{m'(\varphi_0)}{m_p} = -\frac{d}{4}$$

Change in the EMRI dynamics universally captured by the scalar charge

$$G_{\mu\nu} = T_{\mu\nu}^p = 8\pi m_p \int \frac{\delta^{(4)}(x - y_p(\lambda))}{\sqrt{-g}} \frac{dy_p^\alpha}{d\lambda} \frac{dy_p^\beta}{d\lambda} d\lambda$$

$$\square\varphi = -4\pi d m_p \int \frac{\delta^{(4)}(x - y_p(\lambda))}{\sqrt{-g}} d\lambda$$

The perturbation scheme

EMRI small mass ratio naturally leads to use relativistic perturbation theory to describe their evolution

- *Consider linear perturbations of a Schwarzschild background induced by the small body (same for Kerr)*

$$\begin{array}{c} \text{grav-sector} \\ g_{\alpha\beta} = g_{\alpha\beta}^0 + h_{\alpha\beta} \end{array}$$

$$\begin{array}{c} \text{scal-sector} \\ \varphi = \varphi_0 + \varphi_1 \end{array}$$

- *Decompose $h_{\alpha\beta}$ and φ_1 in tensor and scalar spherical harmonics*
- *For the scalar field*

$$\varphi_1(t, r, \theta, \phi) = \sum_{\ell=0}^{\infty} \sum_{m=-\ell}^{\ell} \frac{\delta\varphi_{\ell m}(t, r)}{r} Y_{\ell m}(\theta, \phi)$$

- *Go to the Fourier space, replace into the field's equation and solve for $\delta\varphi_{\ell m}$*

The perturbation scheme

For the gravitational sector

$$h_{\alpha\beta} = h_{\alpha\beta}^{\text{pol}} + h_{\alpha\beta}^{\text{ax}}$$

$(-1)^\ell$ ← → $(-1)^{\ell+1}$

$$\mathbf{h} = \sum_{\ell=0}^{\infty} \sum_{m=-\ell}^{\ell} \left[\mathcal{A}_{\ell m}^{(0)} \mathbf{a}_{\ell m}^{(0)} + \mathcal{A}_{\ell m}^{(1)} \mathbf{a}_{\ell m}^{(1)} + \mathcal{A}_{\ell m} \mathbf{a}_{\ell m} + \mathcal{B}_{\ell m}^{(0)} \mathbf{b}_{\ell m}^{(0)} + \mathcal{B}_{\ell m} \mathbf{b}_{\ell m} + \mathcal{Q}_{\ell m}^{(0)} \mathbf{c}_{\ell m}^{(0)} + \mathcal{Q}_{\ell m} \mathbf{c}_{\ell m} \right. \\ \left. + \mathcal{D}_{\ell m} \mathbf{d}_{\ell m} + \mathcal{G}_{\ell m} \mathbf{g}_{\ell m} + \mathcal{F}_{\ell m} \mathbf{f}_{\ell m} \right]$$

$$\mathbf{b}_{\ell m} = \frac{n_{\ell r}}{\sqrt{2}} \begin{pmatrix} 0 & 0 & 0 & 0 \\ 0 & 0 & Y_{,\theta}^{\ell m} & Y_{,\phi}^{\ell m} \\ 0 & Y_{,\theta}^{\ell m} & 0 & 0 \\ 0 & Y_{,\phi}^{\ell m} & 0 & 0 \end{pmatrix}$$

○ 7 **polar** components + 3 **axial** harmonics

○ For a spherically symmetric background the 2 families decouple

○ In the Regge-Wheeler-Zerilli gauge the components reduce to **1** axial and **1** polar functions

Regge & Wheeler, PRD 108, 1063 (1957)
Zerilli, PRD 2, 2141 (1970)

The wave equations

We have 3 master equations for 3 perturbations

$$e^{-\lambda} = 1 - 2M/r$$

$$\Lambda = \ell(\ell + 1)/2 - 1$$

$$\frac{d^2 R_{\ell m}}{dr_*^2} + \left[\omega^2 - e^{-\lambda} \left(\frac{\ell(\ell + 1)}{r^2} - \frac{6M}{r^3} \right) \right] R_{\ell m} = J_{\text{ax}}$$

GR

Regge-Wheeler

$$\frac{d^2 Z_{\ell m}}{dr_*^2} + \left[\omega^2 - \frac{18M^3 + 18M^2 r \Lambda + 6M r^2 \Lambda^2 + 2r^3 \Lambda^2 (1 + \Lambda)}{r^3 (3M + r \Lambda)} \right] Z_{\ell m} = J_{\text{pol}}$$

Zerilli

$$\frac{d^2 \delta\varphi_{\ell m}}{dr_*^2} + \left[\omega^2 - e^{-\lambda} \left(\frac{\ell(\ell + 1)}{r^2} + \frac{2M}{r^3} \right) \right] \delta\varphi_{\ell m} = J_{\varphi}$$

Scalar field

○ For circular equatorial orbits

$$J_{\varphi} = \boxed{-d} n_p \frac{4\pi P_{\ell m}(\frac{\pi}{2})}{r^{3/2} e^{\lambda}} \sqrt{r - 3M} \delta(r - r_p) \delta(\omega - m\omega_p)$$

Overall scale
sets by the charge



orbit's radius

$$\omega_p = (M/r_p^3)^{1/2}$$



The GW energy flux

The full solutions at infinity/horizon are needed to compute the emitted gravitational wave fluxes

$$\dot{E}_{\text{grav}}^{\pm} = \frac{1}{64\pi} \sum_{\ell=2}^{\infty} \sum_{m=-\ell}^{\ell} \frac{(\ell+2)!}{(\ell-2)!} (\omega^2 |Z_{\ell m}^{\pm}|^2 + 4 |R_{\ell m}^{\pm}|^2) \quad \dot{E}_{\text{scal}}^{\pm} = \frac{1}{32\pi} \sum_{\ell=1}^{\infty} \sum_{m=-\ell}^{\ell} \omega^2 |\delta\varphi_{\ell m}^{\pm}|^2$$

○ The total contribution

$$\dot{E} = \dot{E}_{\text{grav}}^{+} + \dot{E}_{\text{grav}}^{-} + \dot{E}_{\text{scal}}^{+} + \dot{E}_{\text{scal}}^{-} = \dot{E}_{\text{GR}} + \delta\dot{E}_d$$

○ The binary accelerates due to the extra leakage of energy given by the scalar field channel

○ $\delta\dot{E}_d$ enters at the **same** order in **q** as the GR leading dissipative contribution

How much dephasing?

Once we have the total flux emitted by the binary we can determine its adiabatic evolution

- For the orbital phase

$$\frac{dr}{dt} = -\dot{E} \frac{dr}{dE_{\text{orb}}} \quad , \quad \frac{d\Phi}{dt} = \omega_p = \pm \frac{M^{1/2}}{r^{3/2} \pm \chi M^{3/2}}$$

- The total phase can be written as

$$\Phi_d(t) \sim \Phi_{\text{GR}}(t) + \delta\Phi_d(t)$$

- Both contributions are of the same order $\mathcal{O}(1/q)$
- The term $q\delta\Phi_d(t)$ depends only on the scalar charge

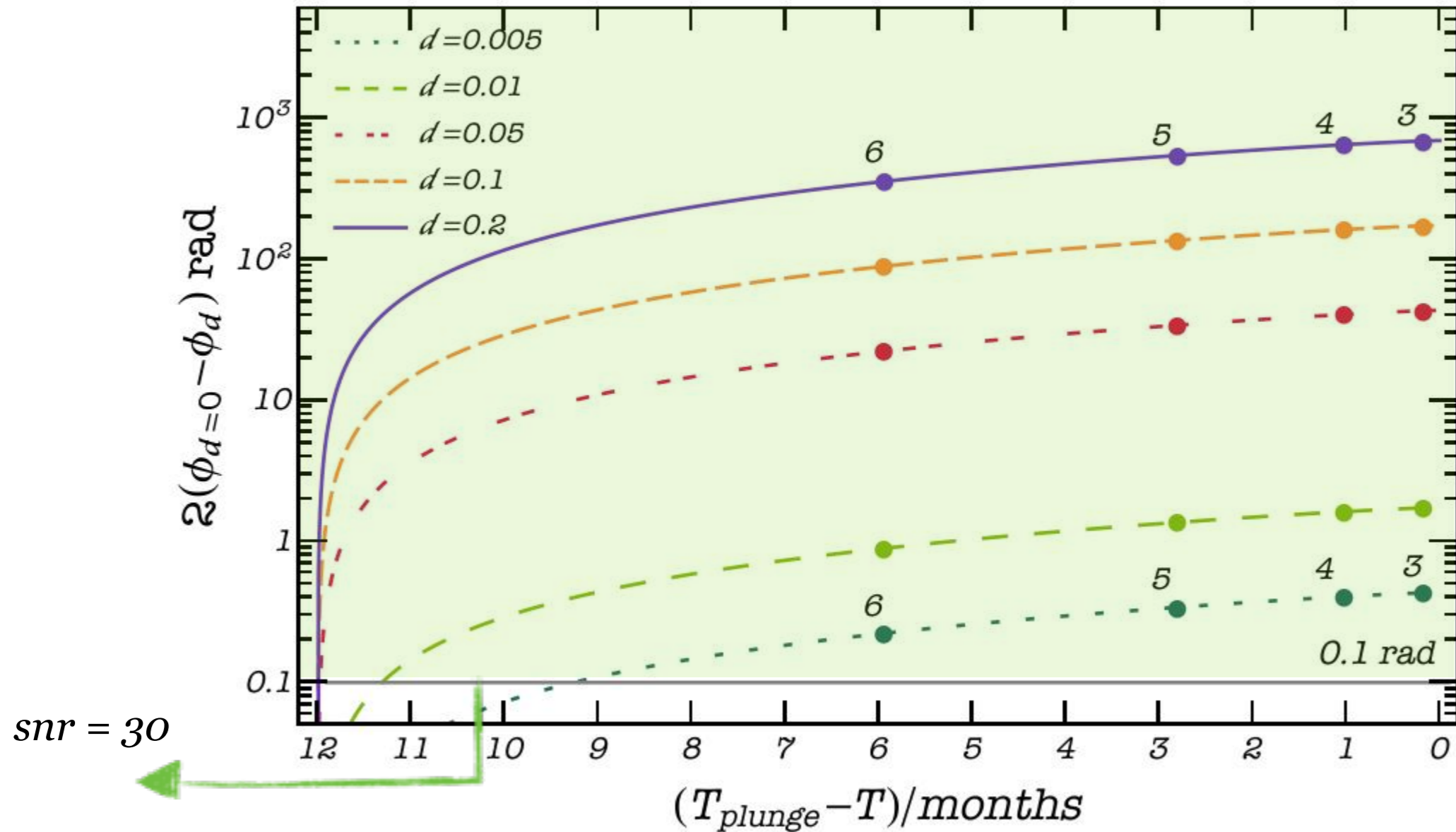
A first assessment of the charge impact is given by studying the **dephasing** induced on the orbital phase

$$\Phi_{d=0}(t) - \Phi_d(t)$$

How much dephasing?

Difference between GR - GR d phase evolution during the inspiral (12 months the plunge)

$$(M, m_p) = (10^6, 10)M_\odot \quad \chi = 0.9$$



○ Potentially able to observe changes induced by scalar charges $d \sim 0.005$

The Waveform

The recipe to generate EMRI waveforms

- *Compute the total energy flux emitted by the binary $\dot{E} = \dot{E}_{\text{GR}} + \delta\dot{E}_d$*
- *The flux drives the binary orbital evolution*

$$\frac{dr(t)}{dt} = -\dot{E} \frac{dr}{dE_{\text{orb}}} \quad , \quad \frac{d\Phi(t)}{dt} = \frac{M^{\frac{1}{2}}}{r_p^{3/2}}$$

- *Build the GW polarizations $h_+[r(t), \Phi(t)]$, $h_\times[r(t), \Phi(t)]$*
- *Given the source localization, construct the strain*

$$h(t) = \frac{\sqrt{3}}{2} [h_+ F_+(\theta, \phi, \psi) + h_\times F_\times(\theta, \phi, \psi)]$$

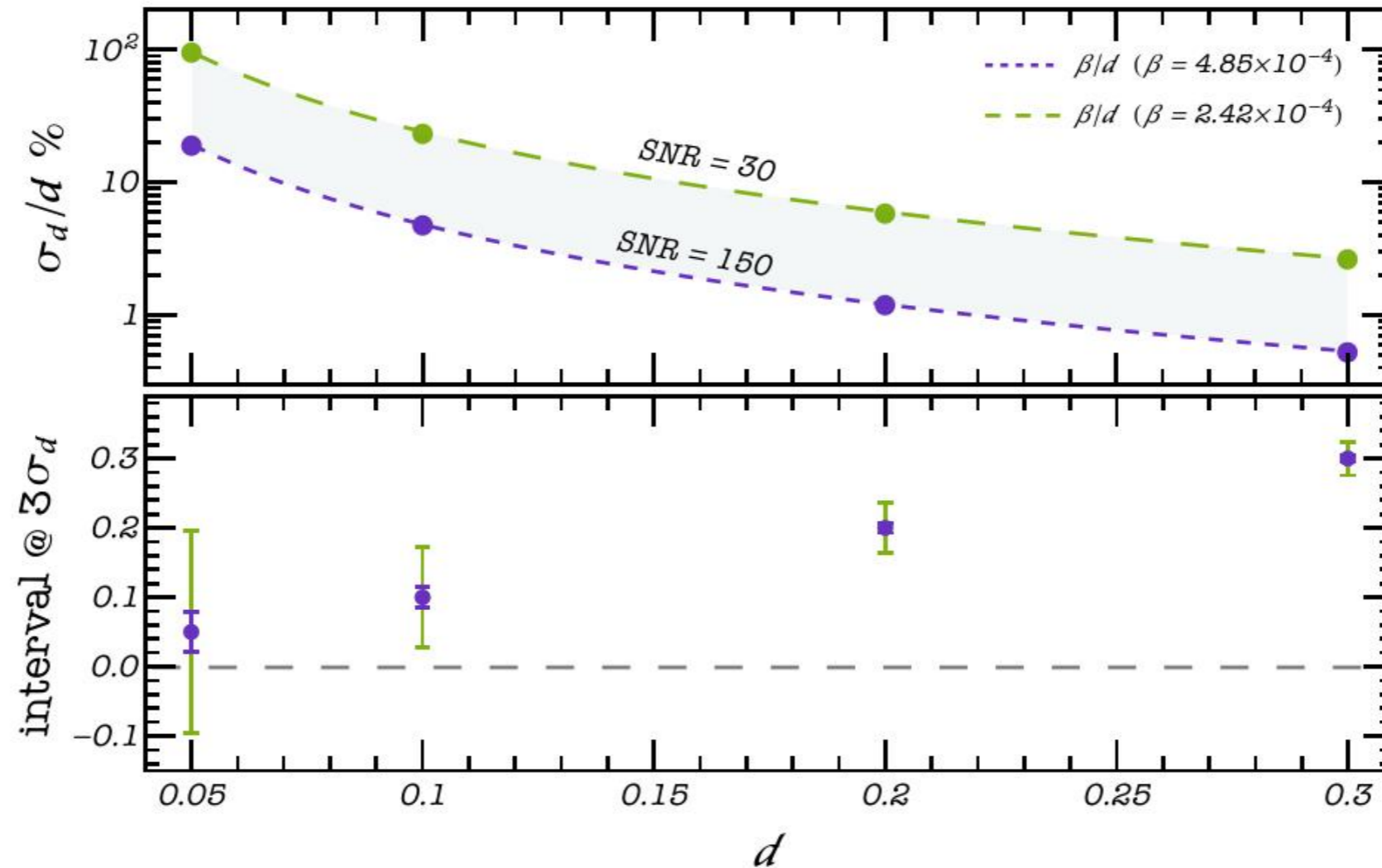
Everything as in GR but $\delta\dot{E}_d$, that only depends on the scalar charge

- *Universal family of waveforms to be tested against GR*

Forecast on LISA bounds

Constraints on the scalar charge for prototype EMRIs with
 $SNR = (30, 150)$

A.M. +, Nature Astronomy 6, 4 464-470 (2022)



- LISA potentially able to measure d with % accuracy and better
- LISA potentially able to constrain $d \sim 0.05$ to be inconsistent with zero @ $3-\sigma$